Black Holes as Bose Einstein Condensates

Tehseen Rug

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2 Quantumness on Macroscopic Scales

3 Connection to Gravity



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Hawking and Black Holes

Hawking: semiclassical treatment of black hole $M_{BH} \rightarrow \infty$, $L_P \rightarrow 0$, r_g fixed \Rightarrow exact thermal spectrum

$$T \sim \frac{1}{r_g}$$
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Question: Is the semiclassical treatment appropriate for real black holes?

Black Holes Quantum Picture

Quantum effects are important, even for macroscopic black holes (Dvali, Gomez).

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$$\alpha = \frac{L_p^2}{\lambda^2} \tag{2}$$

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Regarding black holes as condensates of *N* gravitons $\Rightarrow \alpha = \frac{1}{N}$ Meaning of *N*: single characteristic of black hole / measure of classicality

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$$M_{BH} = \sqrt{N} \frac{\hbar}{L_P}, \ r_g = \sqrt{N} L_P, \ \alpha = \frac{1}{N}$$
 (3)

Resolution of Black Hole Mysteries

Hawking's computation: N $\rightarrow \infty$, L_P \rightarrow 0, r_g fixed.

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Black holes are Bose Einstein Condensates at the critical point of a quantum phase transition!

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Toy Model

Non-relativistic prototype model: bosons on a ring

$$H = \frac{1}{R} \int_0^{2\pi} d\theta \left[-\frac{\hbar^2}{2m} \psi^+ \partial_\theta^2 \psi - \frac{\hbar^2}{2m} \frac{\pi g R}{2} \psi^+ \psi^+ \psi \psi \right]$$
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Performing Bogoliubov with analysis $\hbar = R = 2m = 1$ gives energy spectrum of Bogoliubov modes:

$$\epsilon = \sqrt{k^2(k^2 - gN)}.$$
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- \Rightarrow Quantum phase transition: k = 1 gN = 1.
- \Rightarrow Phase transition is a long wavelength effect.

Measure of Quantumness

Ground state fidelity: $F(gN, gN + \delta) = |\langle 0_{gN} | 0_{gN+\delta} \rangle|$ (Flassig, Pritzel, Wintergerst)

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Figure: Numerical ground state fidelity susceptibility

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Fluctuation entanglement: Neumann entropy of reduced density matrix:

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Fluctuation entanglement: Neumann entropy of reduced density matrix:



Figure: Analytical fluctuation entanglement

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Figure: Numerical fluctuation entanglement

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Black Holes and BEC's

Black Hole - Bose Einstein Condensate correspondence:

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Black Hole - Bose Einstein Condensate correspondence:

self-sustainability \leftrightarrow critical point of quantum phase transition

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Black Hole - Bose Einstein Condensate correspondence:

self-sustainability \leftrightarrow critical point of quantum phase transition Bekenstein entropy \leftrightarrow quantum degeneracy of the BEC at the critical point Hawking radiation \leftrightarrow quantum depletion of the BEC.

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Peculiarity of Gravity

Black holes: $\alpha N = 1$ fixed \Rightarrow Black holes always at the critical point.

Reason: self-similarity of black hole collapse

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Black holes: $\alpha N = 1$ fixed \Rightarrow Black holes always at the critical point.

Reason: self-similarity of black hole collapse Shrinking of size by one quanta:

$$\Delta E = \frac{\hbar}{\sqrt{N}L_P}.$$
(6)

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Same energy needed to leak out one quanta.

- Application of the picture to dS/AdS spaces

$$N = \frac{R^{D-2}}{L_D^{D-2}}$$
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 \rightarrow Connection to AdS/CFT

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- Study higher-dimensional systems

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- Application to Classicalization

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Thank You for Your Attention

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