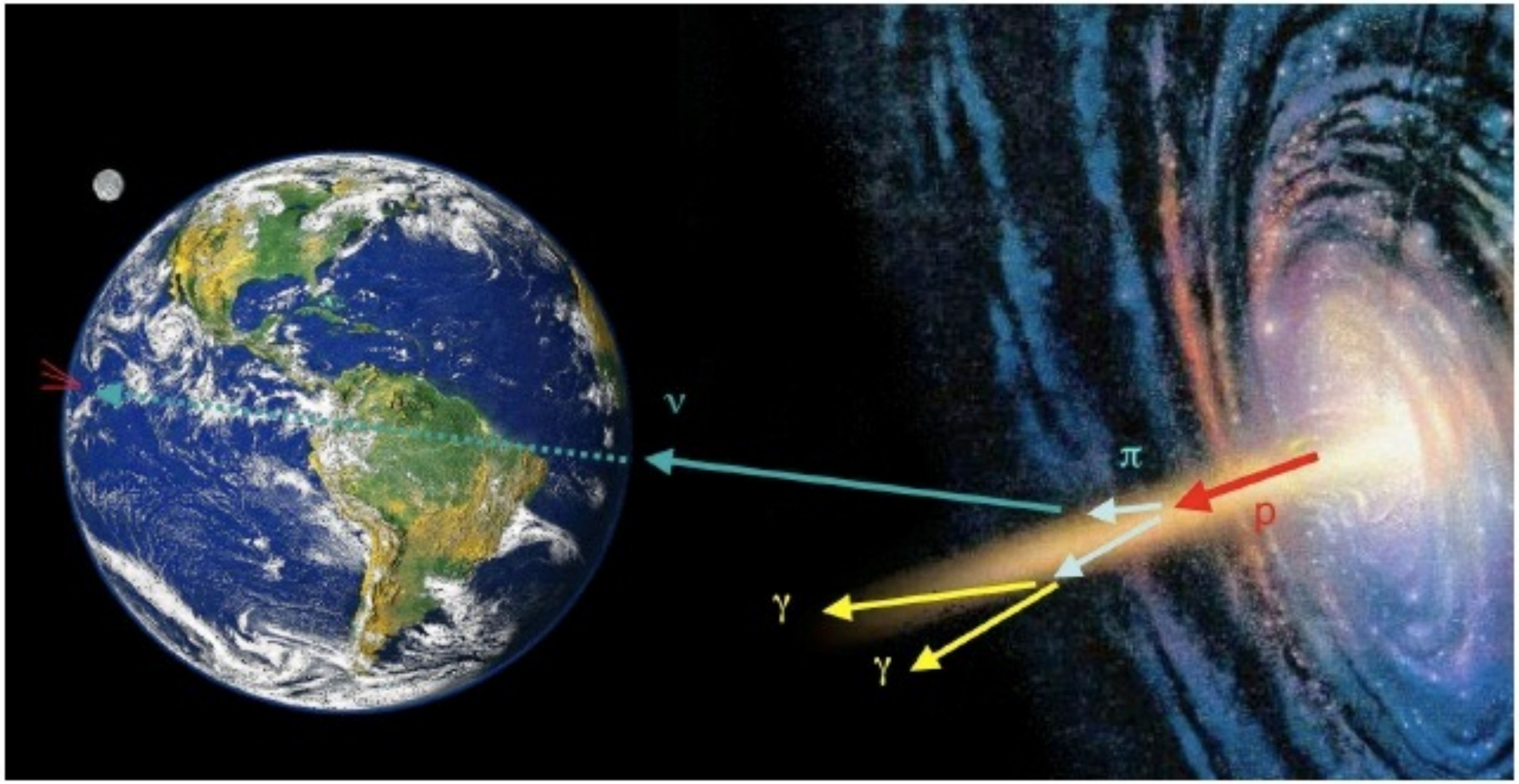
Teilchenphysik mit kosmischen und mit erdgebundenen Beschleunigern



02. Cosmic Accelerators

08.05.2017

Max-Planck-Institut für Physik
(Werner-Heisenberg-Institut)

Prof. Dr. Siegfried Bethke
Dr. Frank Simon

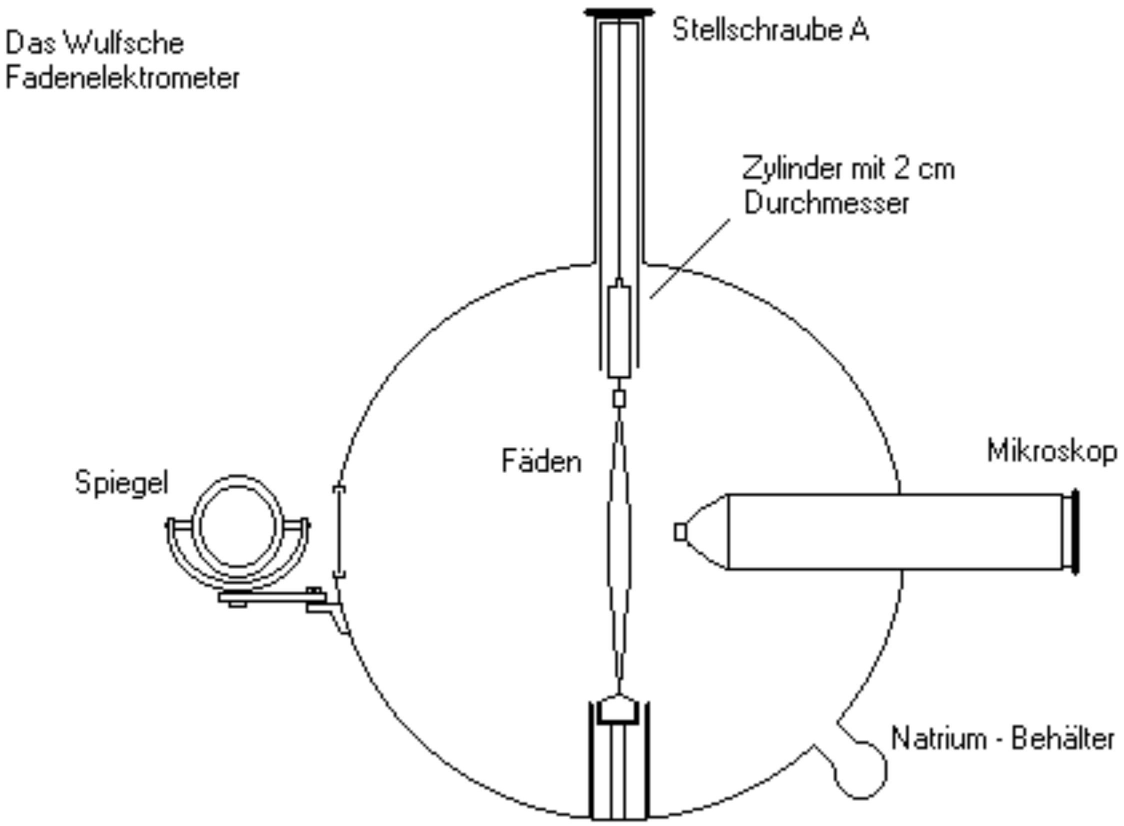
Cosmic Rays: Discovery

- Discovered by Victor Hess 1912
  - Nobel Prize in physics 1936

- Observation on balloon flights with electroscopes:
  - Rate of discharge reduces with increasing altitude, up to an altitude of 1000 m
  - Above this a strong increase of the discharge rate is observed, at 5000 m it is several times higher than the rate at ground level

Cosmic Rays: Discovery

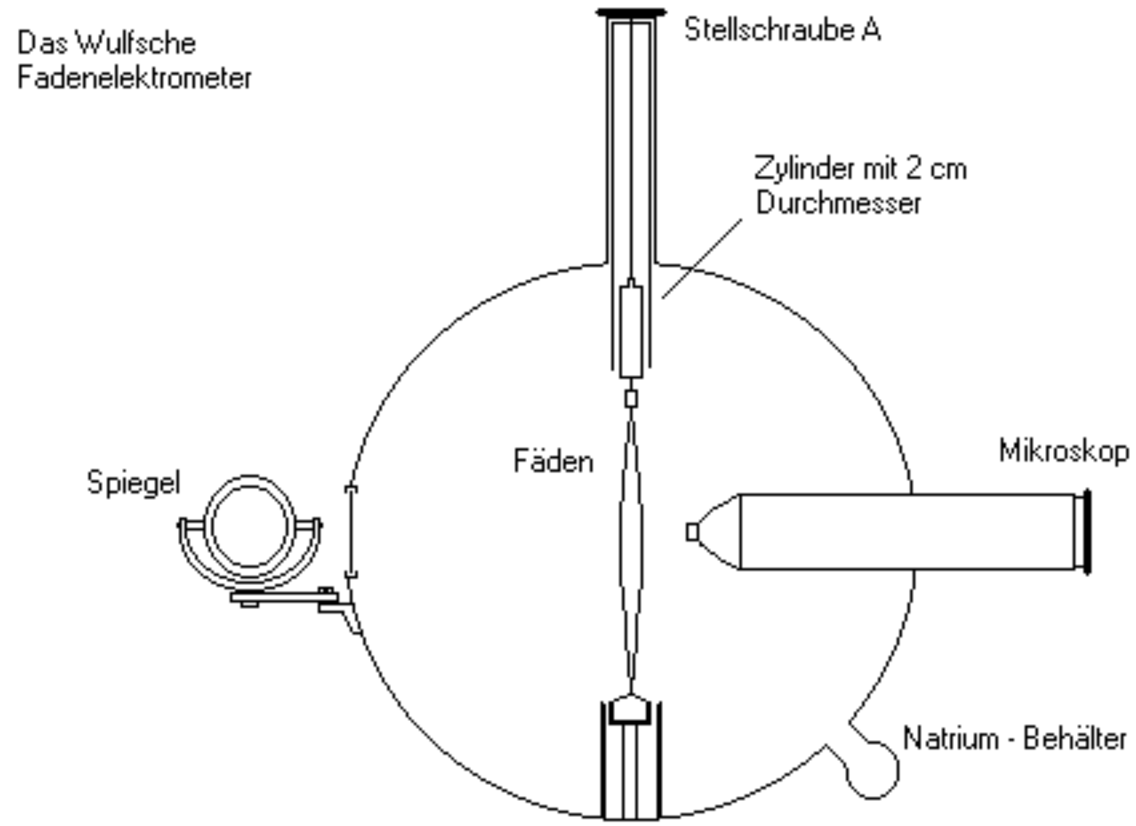
- The experimental method:
    - Electrometer, the distance of (electrostatically charged) strings gives the amount of charge on the strings
    - Discharge via ionising radiation



G. Federmann, Diplomarbeit, U. Wien, 2002

Cosmic Rays: Discovery

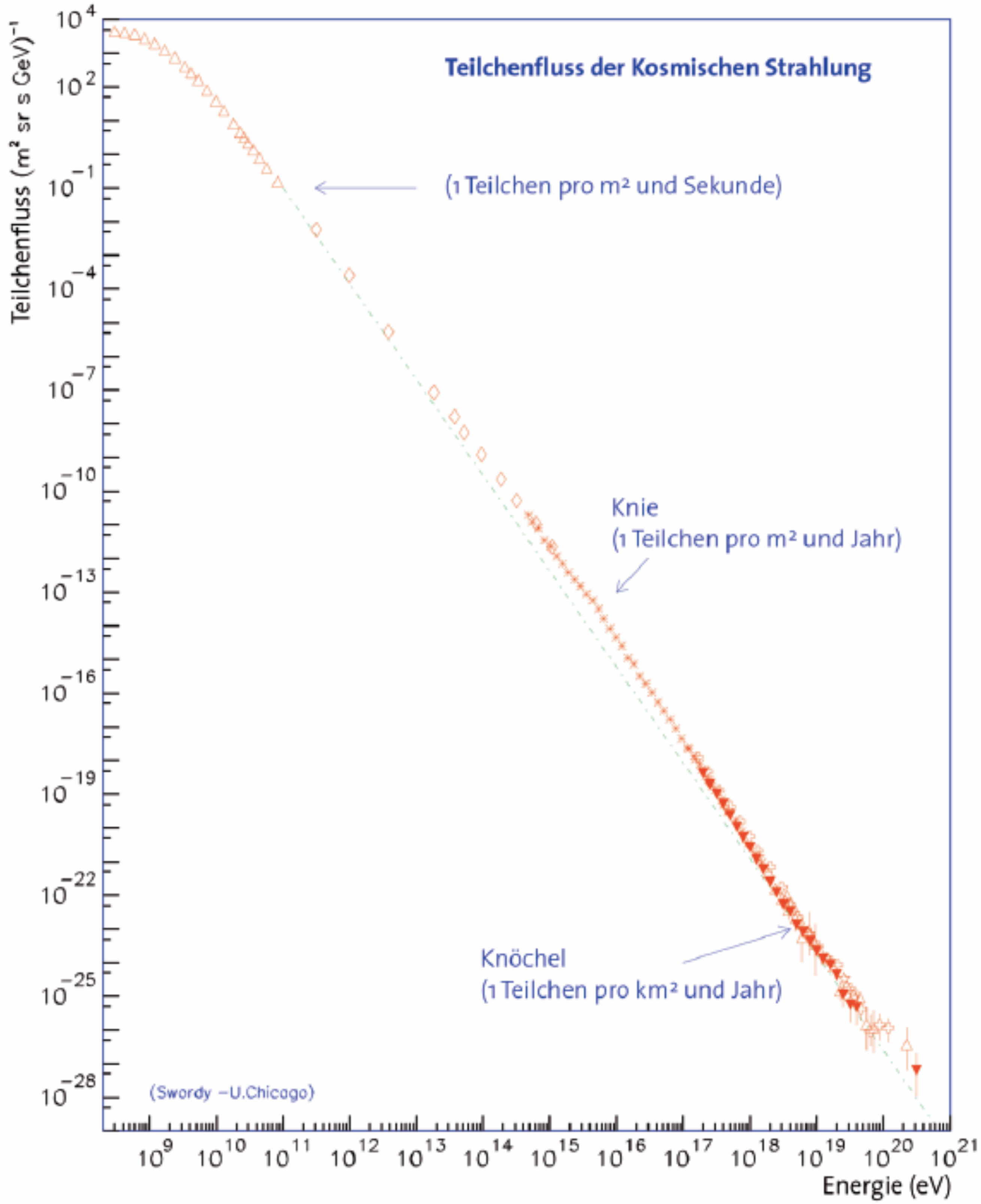
- The experimental method:
  - Electrometer, the distance of (electrostatically charged) strings gives the amount of charge on the strings
  - Discharge via ionising radiation



G. Federmann, Diplomarbeit, U. Wien, 2002

- Interpretation of the observation
  - Reduction of ambient radioactivity with increasing altitude (less radio nuclides, such as Radon)
  - The increase of radiation at high altitudes has to be due to extraterrestrial sources ➪ **"Höhenstrahlung"**

Cosmic Rays: Spectrum



- Extends over many orders of magnitude in energy and flux:
  - GeV (10^9 eV) - ZeV (10^{21})
  - >1 $cm^{-2}s^{-1}$ - < 1 km^{-2} per century

- Follows a power law:

$$\frac{dN}{dE} \propto E^{-\gamma}$$

- $\gamma \sim 2.7$ $E < 10^{15}$ eV
- $\gamma \sim 3.0$ 10^{15} eV < $E < 10^{18}$ eV
- $\gamma \sim 2.7$ 10^{18} eV < E

Energy Density of Cosmic Rays

- Differential flux on earth
  (parametrisation valid from ~ GeV to ~100 TeV):

$$\frac{dN}{dE} \approx \left(\frac{E}{\mathrm{GeV}}\right)^{-2.7} \frac{\mathrm{particles}}{m^2\, sr\, s\, \mathrm{GeV}}$$

Energy Density of Cosmic Rays

- Differential flux on earth
(parametrisation valid from ~ GeV to ~100 TeV):

$$\frac{dN}{dE} \approx \left(\frac{E}{\mathrm{GeV}}\right)^{-2.7} \frac{\text{particles}}{m^2\, sr\, s\, \mathrm{GeV}}$$

- Connection of flux and particle density - assumption: particles travel at speed of light:

$$\text{flux} = \frac{1}{4\pi} \times \text{particle density} \times c$$

Energy Density of Cosmic Rays

- Differential flux on earth
  (parametrisation valid from ~ GeV to ~100 TeV):

$$\frac{dN}{dE} \approx \left(\frac{E}{\mathrm{GeV}}\right)^{-2.7} \frac{\mathrm{particles}}{m^2\, sr\, s\, \mathrm{GeV}}$$

- Connection of flux and particle density - assumption: particles travel at speed of light:

$$\mathrm{flux} = \frac{1}{4\pi} \times \mathrm{particle\ density} \times c$$

- Energy density:

$$\rho_E = \frac{4\pi}{c} \int E \frac{dN}{dE} \approx 1 \frac{eV}{cm^3}$$

Cosmic Rays: Power

- Rough estimate of the energy of cosmic rays within the Milky Way

$$V = \pi R^2 d \;\sim\; \pi (15\,kpc)^2\,(200\,pc) \sim\; 4 \times 10^{60}\,m^3$$

$$\Rightarrow\; E_{Strahlung} \;\sim\; 4 \times 10^{66} eV \sim\; 6.4 \times 10^{47}\,J$$

Cosmic Rays: Power

- Rough estimate of the energy of cosmic rays within the Milky Way

$$V = \pi R^2 d \;\sim\; \pi (15\,kpc)^2\,(200\,pc) \sim\; 4 \times 10^{60}\,m^3$$

$$\Rightarrow\; E_{Strahlung} \;\sim\; 4 \times 10^{66} eV \sim\; 6.4 \times 10^{47}\,J$$

- Assumption: Particles stay within the galaxy for a few million years

$$P_{Strahlung} \sim\; 7 \times 10^{33} W$$

Cosmic Rays: Power

- Rough estimate of the energy of cosmic rays within the Milky Way

$$V = \pi R^2 d \;\sim\; \pi(15\,kpc)^2\,(200\,pc) \sim\; 4 \times 10^{60}\,m^3$$

$$\Rightarrow\; E_{Strahlung} \;\sim\; 4 \times 10^{66} eV \sim\; 6.4 \times 10^{47}\,J$$

- Assumption: Particles stay within the galaxy for a few million years

$$P_{Strahlung} \sim\; 7 \times 10^{33} W$$

- Highly energetic events in the galaxy: supernova explosions

$$E_{SN} \sim\; 10^{44}\,J$$

Cosmic Rays: Power

- Rough estimate of the energy of cosmic rays within the Milky Way

$$V = \pi R^2 d \ \sim \ \pi (15\, kpc)^2\, (200\, pc) \sim \ 4 \times 10^{60}\, m^3$$

$$\Rightarrow \ E_{Strahlung} \ \sim \ 4 \times 10^{66} eV \sim \ 6.4 \times 10^{47}\, J$$

- Assumption: Particles stay within the galaxy for a few million years

$$P_{Strahlung} \sim \ 7 \times 10^{33} W$$

- Highly energetic events in the galaxy: supernova explosions

$$E_{SN} \sim \ 10^{44}\, J$$

- About one supernova every 30 years

$$P_{SN} \sim \ 10^{35}\, W$$

Cosmic Rays: Power

- Rough estimate of the energy of cosmic rays within the Milky Way

$$V = \pi R^2 d \ \sim \ \pi (15\, kpc)^2\, (200\, pc) \sim \ 4 \times 10^{60}\, m^3$$

$$\Rightarrow \ E_{Strahlung} \ \sim \ 4 \times 10^{66} eV \sim \ 6.4 \times 10^{47}\, J$$

- Assumption: Particles stay within the galaxy for a few million years

$$P_{Strahlung} \sim \ 7 \times 10^{33} W$$

- Highly energetic events in the galaxy: supernova explosions

$$E_{SN} \sim \ 10^{44}\, J$$

- About one supernova every 30 years

$$P_{SN} \sim \ 10^{35}\, W$$

‣ SNe could be the accelerators, would require ~10% efficiency

Cosmic Rays: Power

- Rough estimate of the energy of cosmic rays within the Milky Way

$$V = \pi R^2 d \;\sim\; \pi (15\,kpc)^2\,(200\,pc) \sim\; 4 \times 10^{60}\,m^3$$

$$\Rightarrow\; E_{Strahlung} \;\sim\; 4 \times 10^{66} eV \sim\; 6.4 \times 10^{47}\,J$$

- Assumption: Particles stay within the galaxy for a few million years

$$P_{Strahlung} \sim\; 7 \times 10^{33} W$$

- Highly energetic events in the galaxy: supernova explosions

$$E_{SN} \sim\; 10^{44}\,J$$

- About one supernova every 30 years

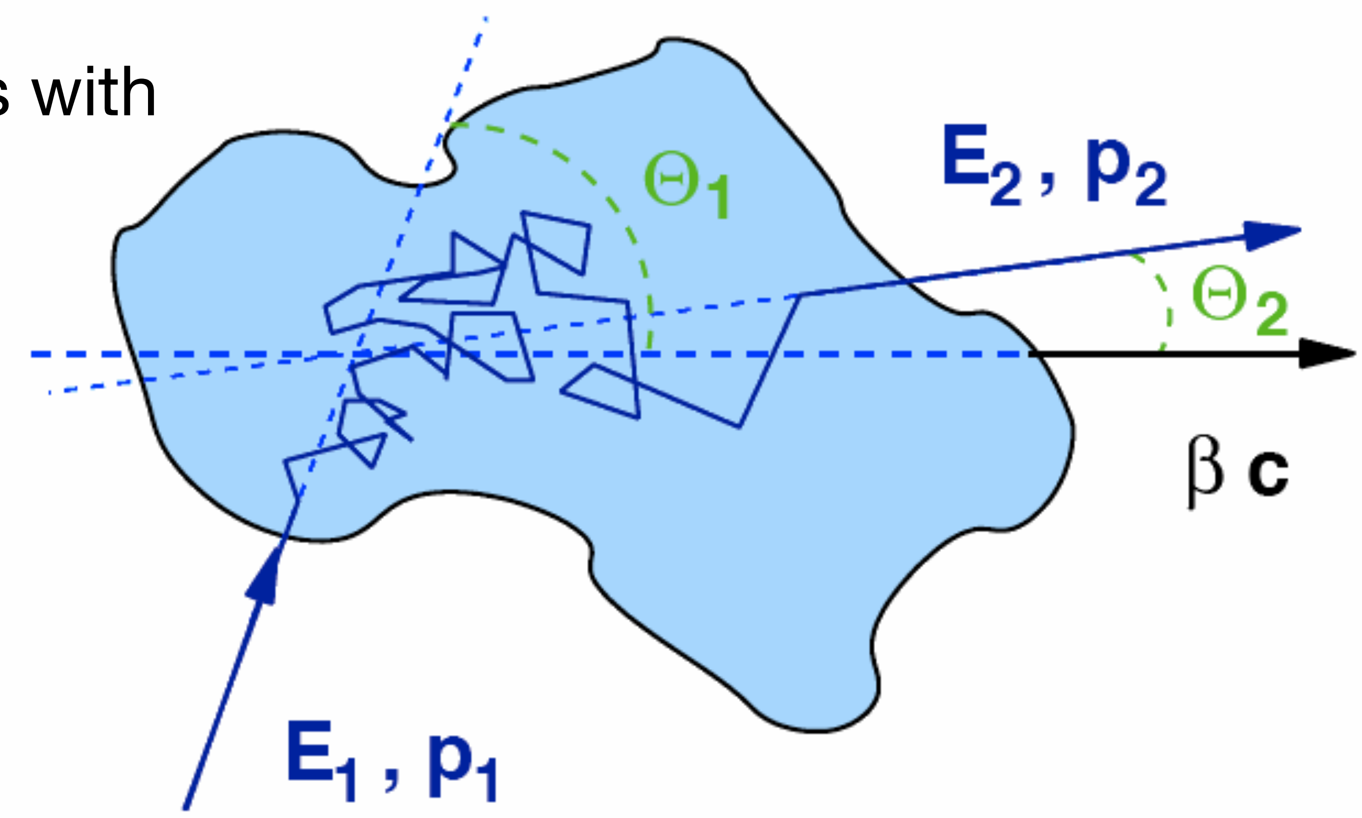
$$P_{SN} \sim\; 10^{35}\,W$$

As comparison:
therm. power of the sun ~4x10^{26} W,
Milky Way ~2x10^{11} stars

- SNe could be the accelerators, would require ~10% efficiency

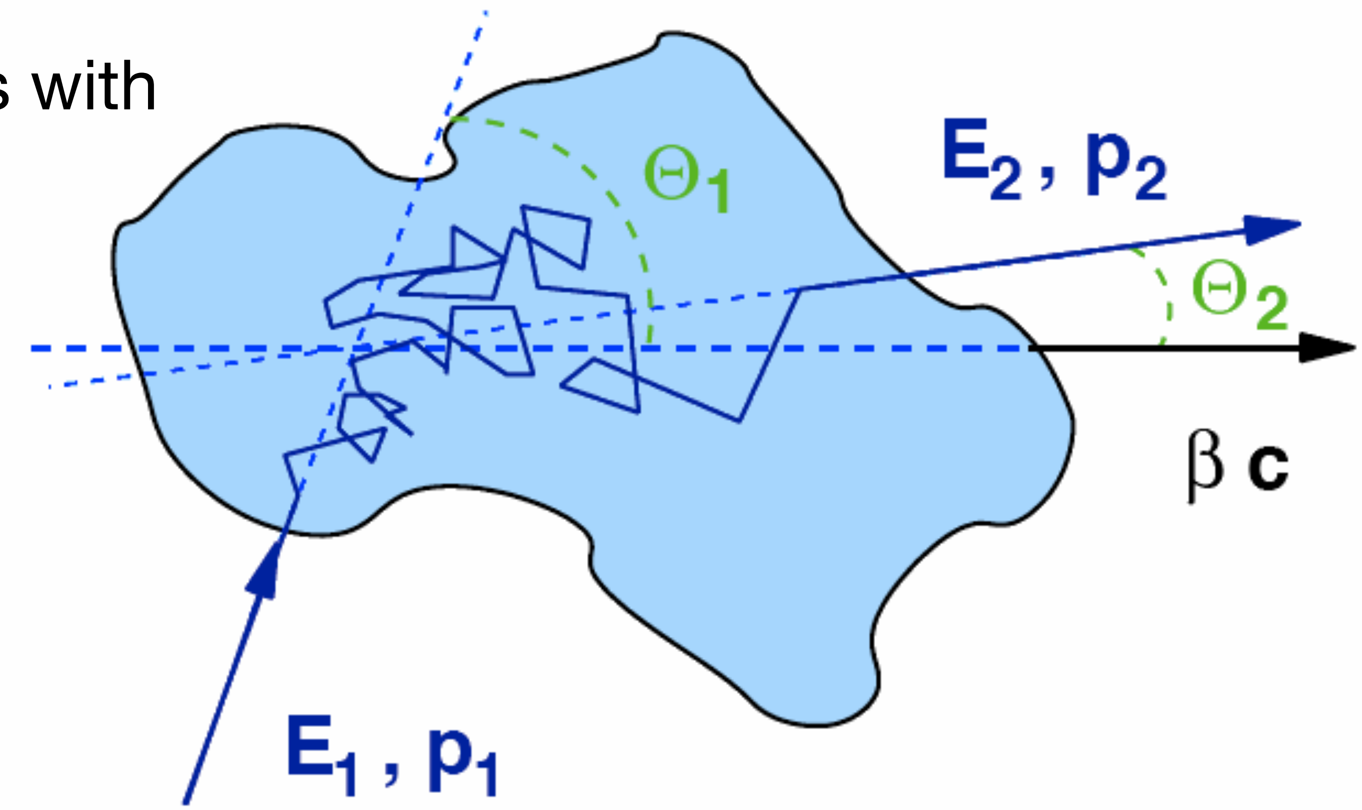
Second Order Fermi Acceleration

- Proposed by Enrico Fermi 1949
- The principle: Collisions of particles with interstellar gas clouds
  - Particle speed ~ c
  - Cloud speed βc
  - entrance and exit angle relative to cloud direction Θ_1, Θ_2



Second Order Fermi Acceleration

- Proposed by Enrico Fermi 1949
- The principle: Collisions of particles with interstellar gas clouds
  - Particle speed ~ c
  - Cloud speed βc
  - entrance and exit angle relative to cloud direction Θ_1, Θ_2

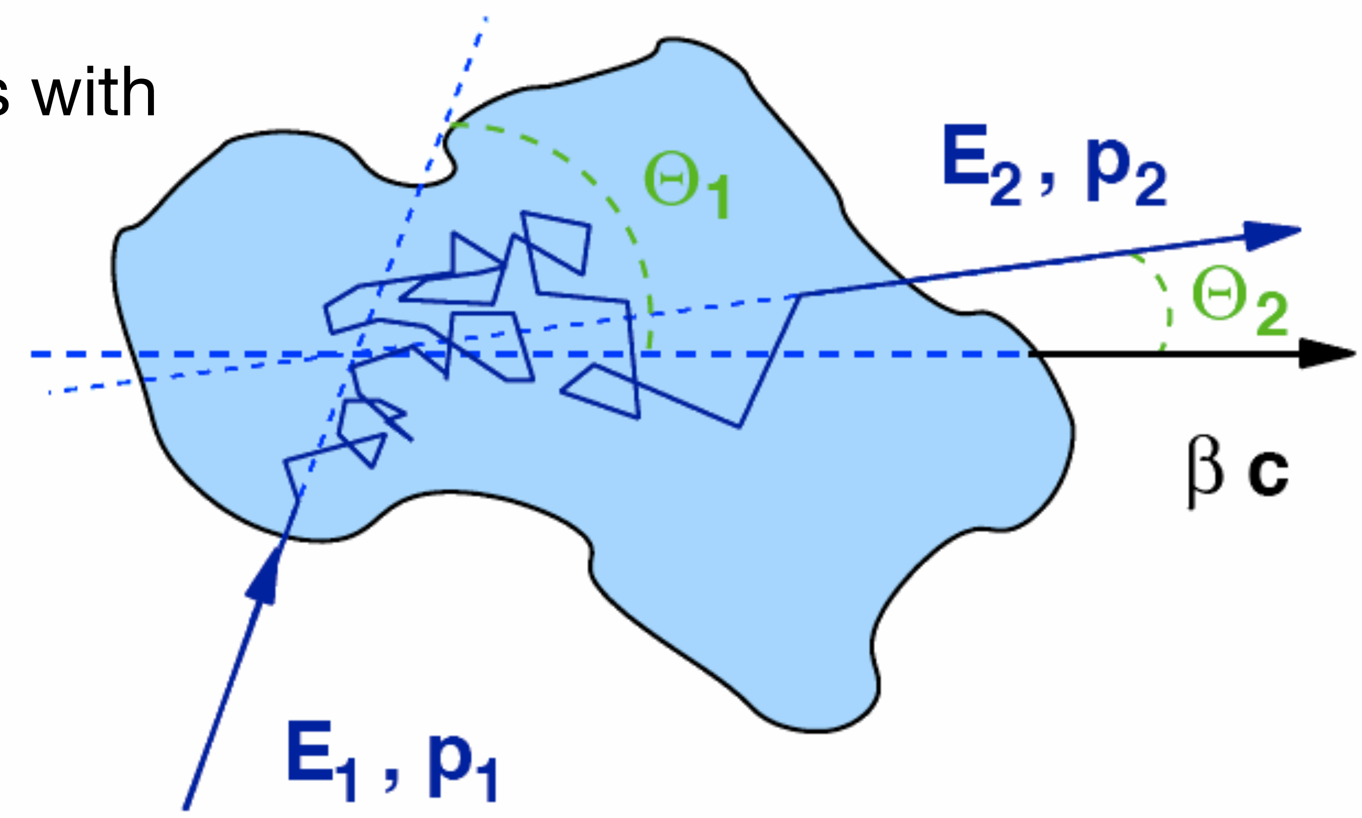


- Boost into cloud rest frame:

$$E_1' = \gamma E_1 - \beta\gamma p_{||} \;=\; \gamma E_1 - \beta\gamma cos\Theta_1\, p \approx \; \gamma E_1(1 - \beta cos\Theta_1)$$

Second Order Fermi Acceleration

- Proposed by Enrico Fermi 1949
- The principle: Collisions of particles with interstellar gas clouds
  - Particle speed ~ c
  - Cloud speed βc
  - entrance and exit angle relative to cloud direction Θ_1, Θ_2



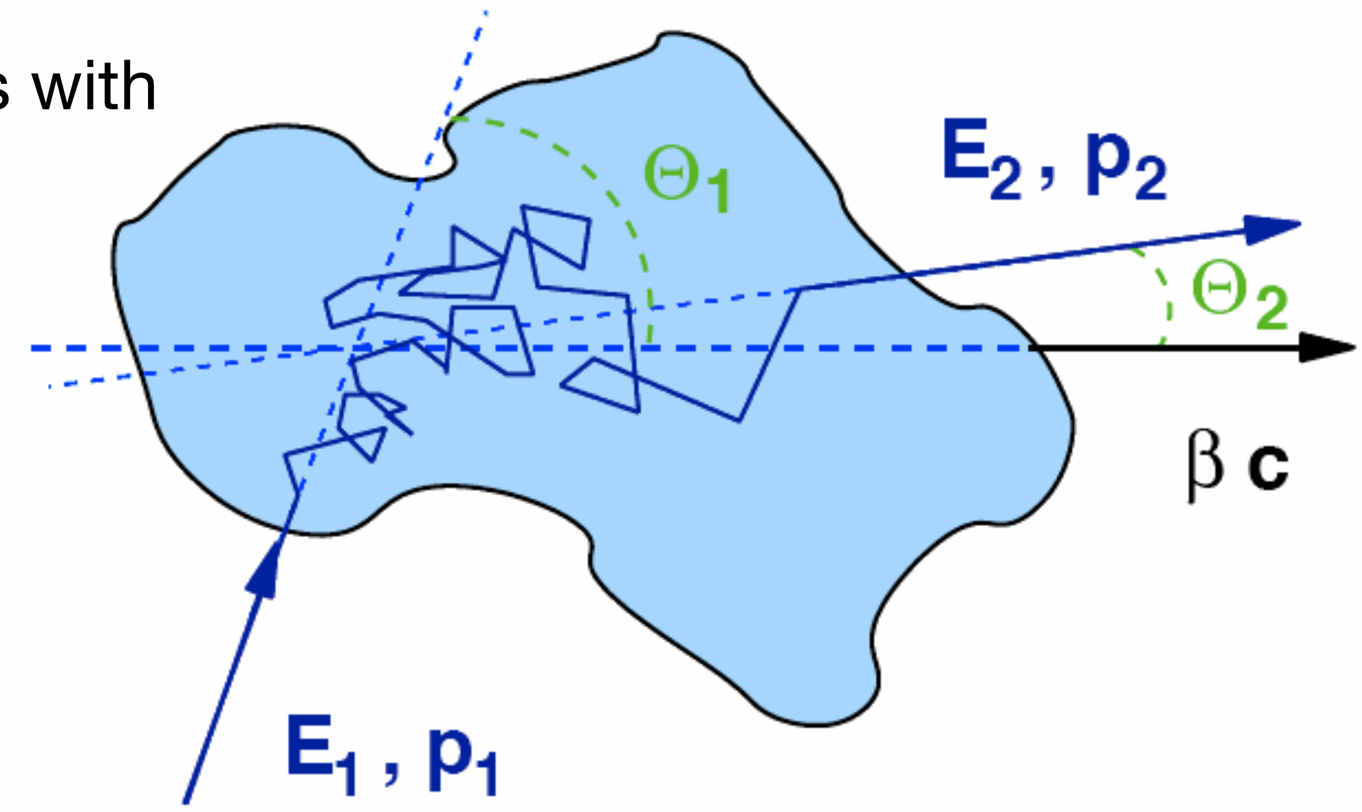
- Boost into cloud rest frame:

$$E'_1 = \gamma E_1 - \beta\gamma p_{||} \;=\; \gamma E_1 - \beta\gamma cos\Theta_1\, p \approx \; \gamma E_1(1 - \beta cos\Theta_1)$$

- Elastic scattering within the cloud: $E'_2 = E'_1$ $\qquad \langle cos\Theta'_2 \rangle = 0$

Second Order Fermi Acceleration

- Proposed by Enrico Fermi 1949
- The principle: Collisions of particles with interstellar gas clouds
  - Particle speed ~ c
  - Cloud speed βc
  - entrance and exit angle relative to cloud direction Θ_1, Θ_2



- Boost into cloud rest frame:

$$E_1' = \gamma E_1 - \beta\gamma p_{||} \; = \; \gamma E_1 - \beta\gamma cos\Theta_1\, p \approx \; \gamma E_1(1 - \beta cos\Theta_1)$$

- Elastic scattering within the cloud: $E_2' = E_1'$ $\qquad \langle cos\Theta_2' \rangle = 0$

- Boost boost back to "Universe" frame: $E_2 = \gamma E_2'(1 + \beta cos\Theta_2')$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

$$\Rightarrow \; E_2 = \gamma(\gamma E_1 - \beta\gamma E_1 cos\Theta_1)(1 + \beta cos\Theta_2')$$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

$$\Rightarrow \quad E_2 = \gamma(\gamma E_1 - \beta\gamma E_1 cos\Theta_1)(1 + \beta cos\Theta_2')$$

$$\Rightarrow \quad \frac{\Delta E}{E} = \gamma^2(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2') - 1$$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

$$\Rightarrow \quad E_2 = \gamma(\gamma E_1 - \beta\gamma E_1 cos\Theta_1)(1 + \beta cos\Theta_2')$$

$$\Rightarrow \quad \frac{\Delta E}{E} = \gamma^2(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2') - 1$$

$$= \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

$$\Rightarrow \quad E_2 = \gamma(\gamma E_1 - \beta\gamma E_1 cos\Theta_1)(1 + \beta cos\Theta_2')$$

$$\Rightarrow \quad \frac{\Delta E}{E} = \gamma^2(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2') - 1$$

$$= \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

- Mean values: $\langle cos\Theta_2' \rangle = 0$
    - Scattering probability depends on relative velocity between particle and cloud

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

$$\Rightarrow \; E_2 = \gamma(\gamma E_1 - \beta\gamma E_1 cos\Theta_1)(1 + \beta cos\Theta_2')$$

$$\Rightarrow \; \frac{\Delta E}{E} = \gamma^2(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2') - 1$$

$$= \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

- Mean values: $\langle cos\Theta_2' \rangle = 0$
    - Scattering probability depends on relative velocity between particle and cloud

$$P \propto (c - c\beta cos\Theta_1) \propto 1 - \beta cos\Theta_1$$

Second Order Fermi Acceleration

- Energy difference $$\frac{\Delta E}{E} = \frac{E_2 - E_1}{E_1} = \frac{E_2}{E_1} - 1$$

$$E_2 = \gamma E_2'(1 + \beta cos\Theta_2') \qquad E_2' = E_1' = \gamma E_1(1 - \beta cos\Theta_1)$$

$$\Rightarrow \ E_2 = \gamma(\gamma E_1 - \beta\gamma E_1 cos\Theta_1)(1 + \beta cos\Theta_2')$$

$$\Rightarrow \ \frac{\Delta E}{E} = \gamma^2(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2') - 1$$

$$= \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

- Mean values: $\langle cos\Theta_2' \rangle = 0$
    - Scattering probability depends on relative velocity between particle and cloud

$$P \propto (c - c\beta cos\Theta_1) \propto 1 - \beta cos\Theta_1$$

$$\Rightarrow \ \frac{dN}{dcos\Theta_1} \propto 1 - \beta cos\Theta_1 \ \Rightarrow \ \langle cos\Theta_1 \rangle = \frac{\int_{-1}^{1} \frac{dN}{dcos\Theta_1} \, cos\Theta_1 \, dcos\Theta_1}{\int_{-1}^{1} \frac{dN}{dcos\Theta_1} \, dcos\Theta_1} = -\frac{\beta}{3}$$

Second Order Fermi Acceleration

- Energy difference: $$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

Second Order Fermi Acceleration

- Energy difference: $$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

  with: $$\langle cos\Theta_1 \rangle = -\frac{\beta}{3} \ , \ \langle cos\Theta_2' \rangle = 0$$

Second Order Fermi Acceleration

- Energy difference: $\dfrac{\Delta E}{E} = \dfrac{(1-\beta cos\Theta_1)(1+\beta cos\Theta_2')}{1-\beta^2} - 1$

  with: $\langle cos\Theta_1 \rangle = -\dfrac{\beta}{3}\ ,\ \langle cos\Theta_2' \rangle = 0$

  $$\Rightarrow \frac{\langle \Delta E \rangle}{E} = \frac{1+\frac{\beta^2}{3}}{1-\beta^2} - 1 \approx \frac{4}{3}\beta^2 \text{ für } \beta \ll 1$$

Second Order Fermi Acceleration

- Energy difference: $$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta'_2)}{1 - \beta^2} - 1$$

with: $$\langle cos\Theta_1 \rangle = -\frac{\beta}{3} \; , \; \langle cos\Theta'_2 \rangle = 0$$

$$\Rightarrow \frac{\langle \Delta E \rangle}{E} = \frac{1 + \frac{\beta^2}{3}}{1 - \beta^2} - 1 \approx \frac{4}{3}\beta^2 \text{ für } \beta \ll 1$$

- Very low acceleration efficiency (proportional to β^2 - second order in β):
  - typical cloud speeds 10^4 m/s => $\beta \sim 3 \times 10^{-5}$
  - mean free path between collisions: ~ 30 pc => every 100 years one collision

Second Order Fermi Acceleration

- Energy difference: $$\frac{\Delta E}{E} = \frac{(1-\beta cos\Theta_1)(1+\beta cos\Theta_2')}{1-\beta^2} - 1$$

with: $$\langle cos\Theta_1 \rangle = -\frac{\beta}{3}\ ,\ \langle cos\Theta_2' \rangle = 0$$

$$\Rightarrow \frac{\langle \Delta E \rangle}{E} = \frac{1+\frac{\beta^2}{3}}{1-\beta^2} - 1 \approx \frac{4}{3}\beta^2 \text{ für } \beta \ll 1$$

- Very low acceleration efficiency (proportional to β^2 - second order in β):
  - typical cloud speeds 10^4 m/s => β ~ 3 x 10^{-5}
  - mean free path between collisions: ~ 30 pc => every 100 years one collision

Energy gain: $$\frac{dE}{dt} = E\,\frac{4\beta^2}{3\tau} \Rightarrow E(t) = E_0\, e^{\frac{4\beta^2}{3\tau}t}$$

Second Order Fermi Acceleration

- Energy difference: $$\frac{\Delta E}{E} = \frac{(1-\beta cos\Theta_1)(1+\beta cos\Theta_2')}{1-\beta^2} - 1$$

with: $$\langle cos\Theta_1 \rangle = -\frac{\beta}{3} \; , \; \langle cos\Theta_2' \rangle = 0$$

$$\Rightarrow \frac{\langle \Delta E \rangle}{E} = \frac{1+\frac{\beta^2}{3}}{1-\beta^2} - 1 \approx \frac{4}{3}\beta^2 \text{ für } \beta \ll 1$$

- Very low acceleration efficiency (proportional to β^2 - second order in β):
  - typical cloud speeds 10^4 m/s => $\beta \sim 3 \times 10^{-5}$
  - mean free path between collisions: ~ 30 pc => every 100 years one collision

Energy gain: $$\frac{dE}{dt} = E\,\frac{4\beta^2}{3\tau} \Rightarrow E(t) = E_0\, e^{\frac{4\beta^2}{3\tau}t}$$

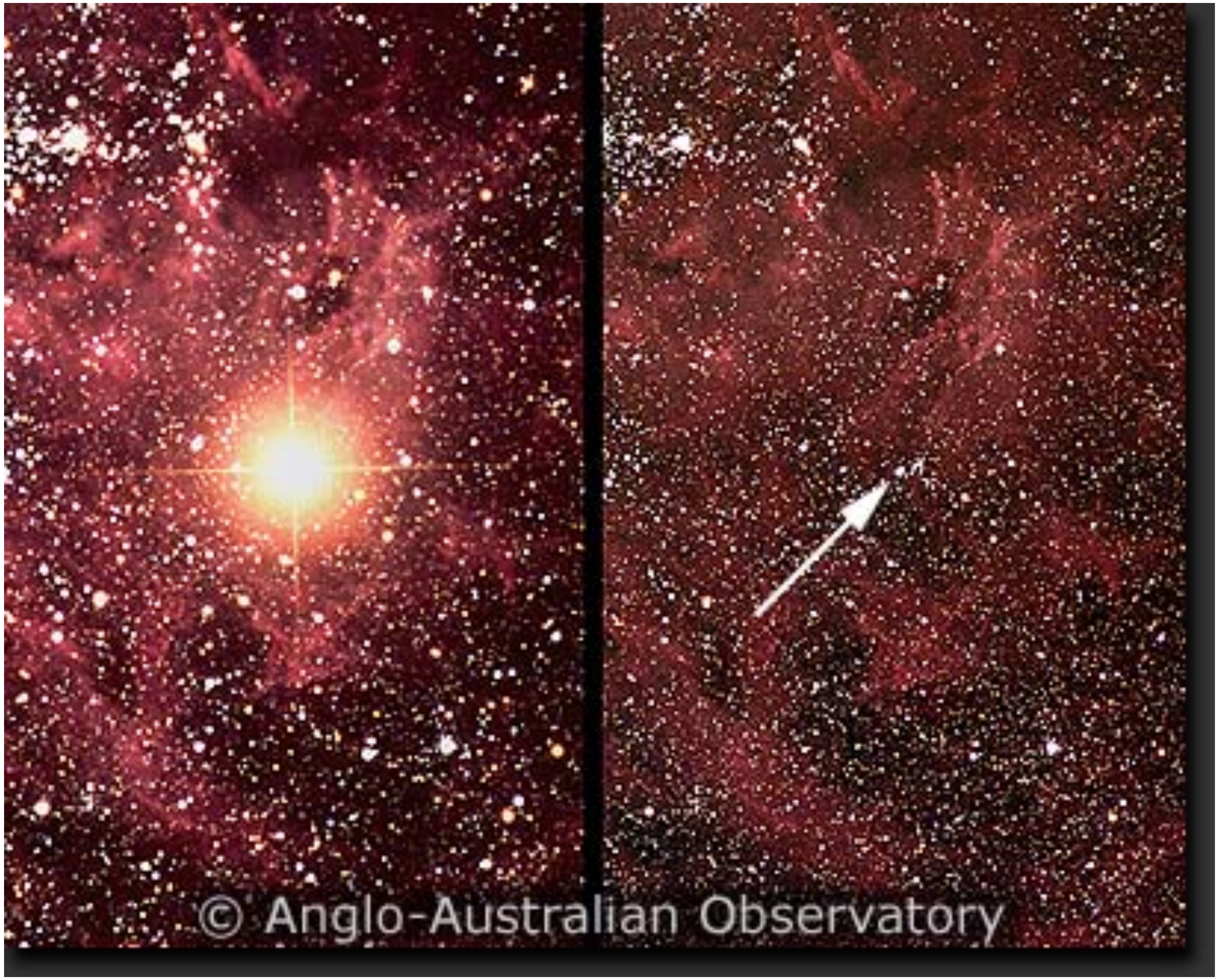
Characteristic time: $$\frac{3\tau}{4\beta^2} \approx 6 \times 10^{10} a$$

Supernovae

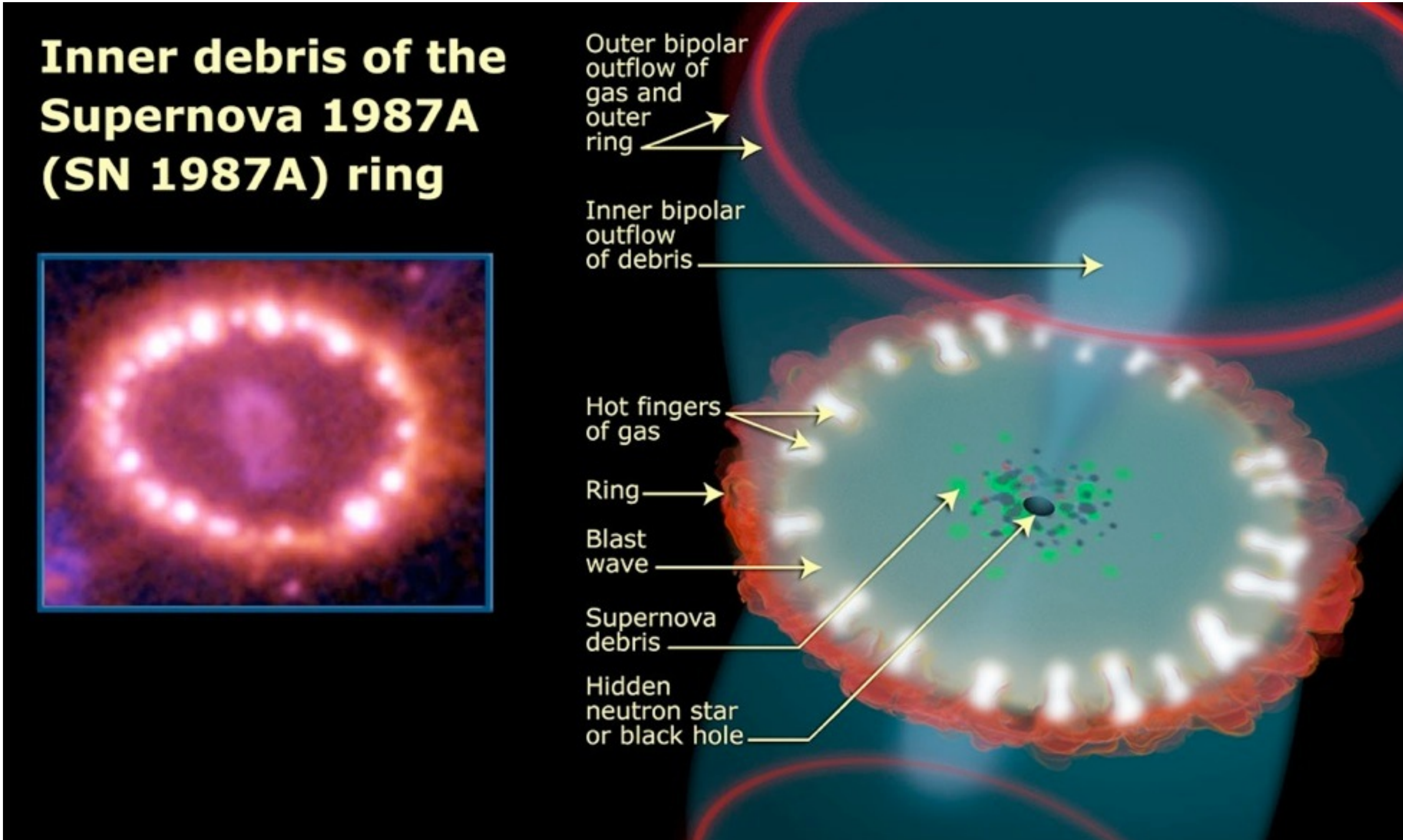
- Classification in to main types
  - SN I: no hydrogen lines in spectrum
    - SN Ia collapse of an accreting white dwarf in a binary star system to a neutron star
  - SN II: hydrogen lines visible
    - Gravitational collapse of a massive star at the end of its life
    - Star burns up to the formation of iron in the core, then no radiation pressure to counter gravitation
      - Atoms are converted to neutrons via electron capture
      - Star collapses with a speed of ~ 0.1 c
      - Matter is reflected at the stable neutron star in the core
      - A shock wave runs outwards
      - An enormous number of neutrinos is produced (~10^{58}), despite their small interaction cross section they further drive the shock wave

Supernova SN1987a

- Supernova explosion 1987 in the great Magellanic Cloud (small partner galaxy of the Milky Way)



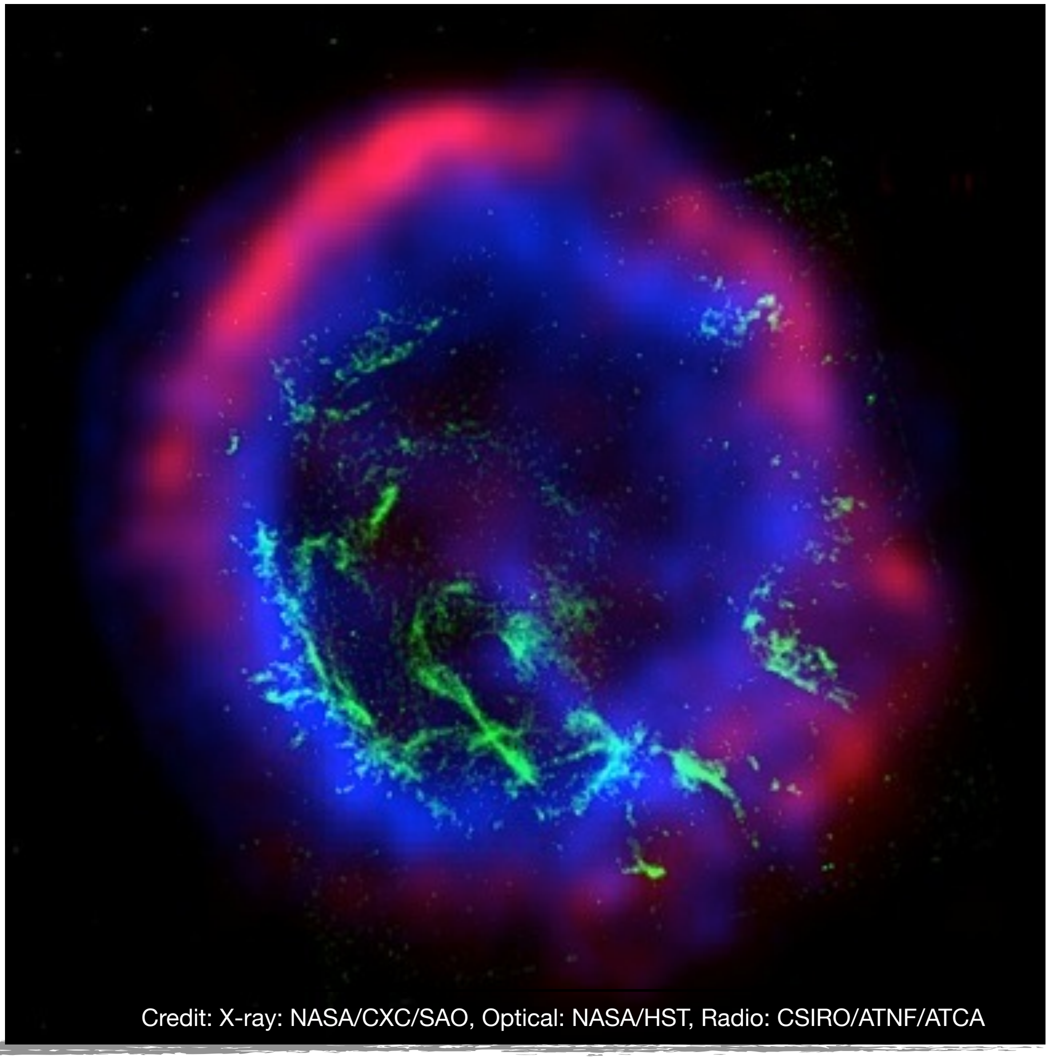
Supernova SN1987a



Credit:NASA, ESA, and A, Feild (STScI)

First Order Fermi Acceleration

- Extension of Fermi acceleration concept to supernova shocks
- Formation of the shock wave:
  - SN ejects large amounts of material (several solar masses) with high velocity into the interstellar medium
  - $v_{material} >> v_{sound(ISM)}$, $v_{material} \sim 10^7$ m/s , $v_{sound(ISM)} \sim 10^4$ m/s
  - ▸ Since the matter is much faster than the speed of sound a shock front develops

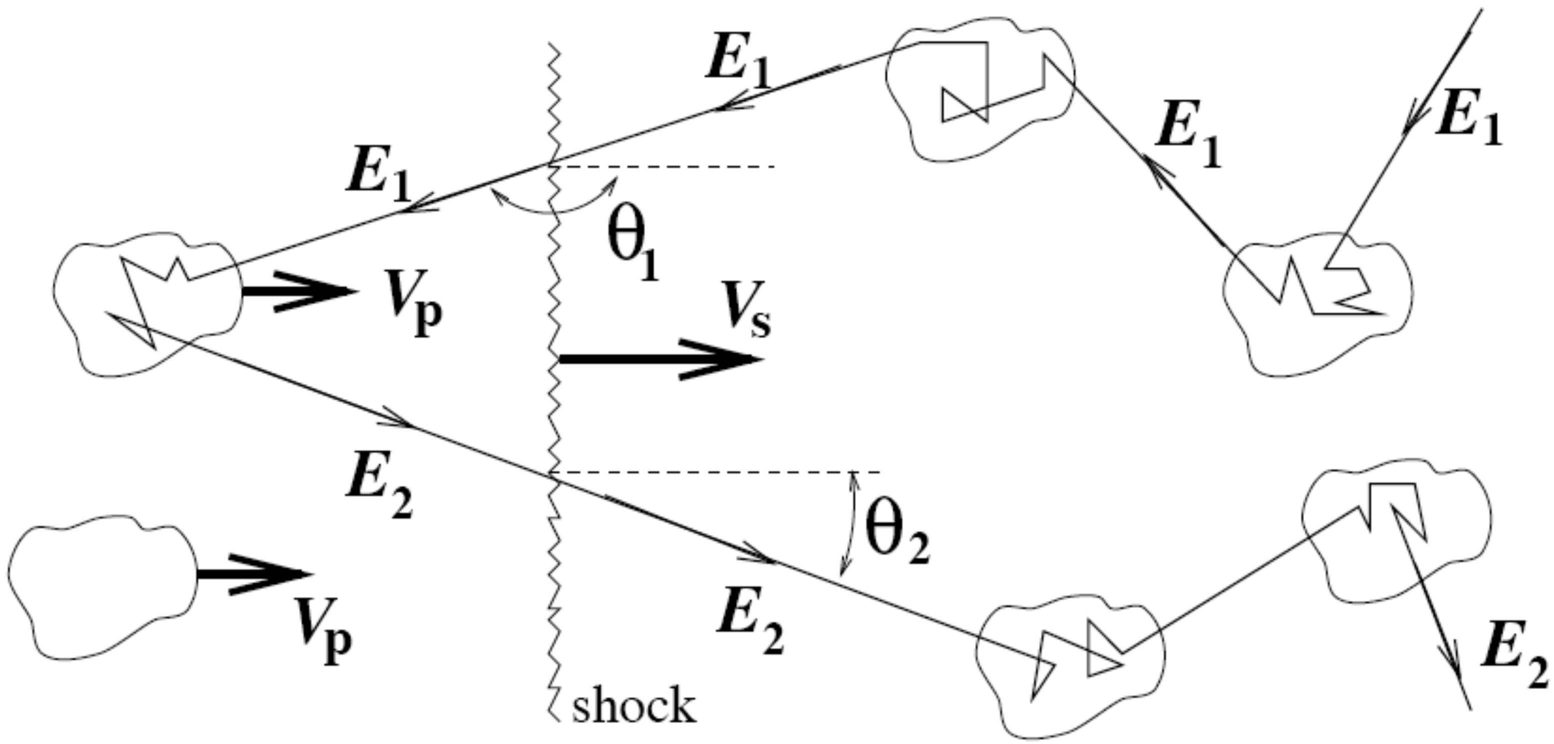


  - The shock wave propagates in the ionized plasma of the ISM (single atoms, specific heat 5/3)
  - Hydrodynamics can show that the speed of the shock wave is:
    $v_{shock}/v_{material} \sim 4/3$

First Order Fermi Acceleration

- Particle acceleration by repeated crossing of shock fronts

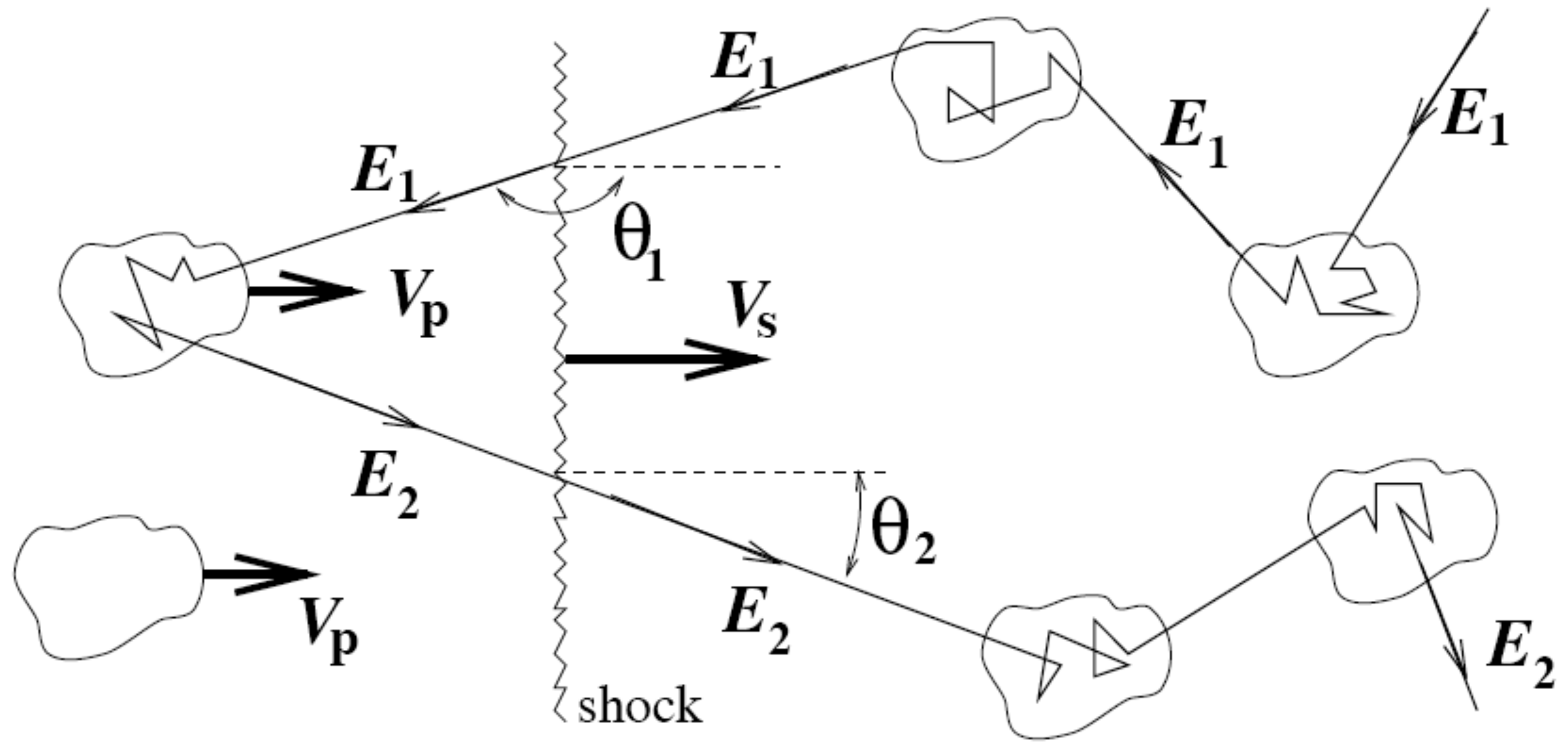
- As for second order Fermi Acceleration the particles are scattered by magnetic field inhomogeneities / turbulences
  - Key “feature”: behind the shock, these turbulences move with the speed of the ejected matter



First Order Fermi Acceleration

- Particle acceleration by repeated crossing of shock fronts

- As for second order Fermi Acceleration the particles are scattered by magnetic field inhomogeneities / turbulences
  - Key “feature”: behind the shock, these turbulences move with the speed of the ejected matter



- Considerations concerning incidence angles of particles:
  - Shock front introduces directional motion - projection of flux onto front

Effective area $A' = -A\,cos\Theta_1$

Flux crossing shock $\frac{dN}{dcos\Theta_1} \propto -cos\Theta_1$

First Order Fermi Acceleration

- Mean value of angles (key point here: Crossing of shock only for $cos\Theta_1 < 0$, other particles do not contribute!):

$$\langle cos\Theta_1 \rangle = \frac{\int_{-1}^{0} \frac{dN}{dcos\Theta_1} cos\Theta_1 \, dcos\Theta_1}{\int_{-1}^{0} \frac{dN}{dcos\Theta_1} \, dcos\Theta_1}$$

$$= \frac{\int_{-1}^{0} -cos^2\Theta_1 \, dcos\Theta_1}{\int_{-1}^{0} -cos\Theta_1 \, dcos\Theta_1} = -\frac{2}{3}$$

analog: $\langle cos\Theta_2 \rangle = \frac{2}{3}$

First Order Fermi Acceleration

- Mean energy change (analogous to second order FA):

$$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

β: Speed of matter behind shock wave

First Order Fermi Acceleration

- Mean energy change (analogous to second order FA):

$$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

β: Speed of matter behind shock wave

$$\frac{\langle \Delta E \rangle}{E} = \frac{(1 + \frac{2}{3}\beta)(1 + \frac{2}{3}\beta)}{1 - \beta^2} - 1$$

First Order Fermi Acceleration

- Mean energy change (analogous to second order FA):

$$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

β: Speed of matter behind shock wave

$$\frac{\langle \Delta E \rangle}{E} = \frac{(1 + \frac{2}{3}\beta)(1 + \frac{2}{3}\beta)}{1 - \beta^2} - 1$$

- for β << 1

$$\frac{\langle \Delta E \rangle}{E} \approx \frac{4}{3}\beta$$

First Order Fermi Acceleration

- Mean energy change (analogous to second order FA):

$$\frac{\Delta E}{E} = \frac{(1 - \beta cos\Theta_1)(1 + \beta cos\Theta_2')}{1 - \beta^2} - 1$$

β: Speed of matter behind shock wave

$$\frac{\langle \Delta E \rangle}{E} = \frac{(1 + \frac{2}{3}\beta)(1 + \frac{2}{3}\beta)}{1 - \beta^2} - 1$$

- for β << 1 $$\frac{\langle \Delta E \rangle}{E} \approx \frac{4}{3}\beta$$

- Substantially more efficient than second order acceleration due to two effects:
  - large velocity differences of material before and after shock front
  - directed motion of shock instead of random drifting
- $\beta \sim 3 \times 10^{-2}$, acceleration linear in β (first order in β)

Energy Spectrum

- Energy gain per cycle

$$\frac{\langle \Delta E \rangle}{E} = \zeta \approx \frac{4}{3}\beta_{Plasma} \approx \beta_{Schock}$$

$$\Rightarrow \quad E = E_0 \left(1 + \zeta\right)^k$$ after k cycles

- Reminder:
  $v_{shock}/v_{material} \sim 4/3$
  for strong shocks

Energy Spectrum

- Energy gain per cycle

$$\frac{\langle \Delta E \rangle}{E} = \zeta \approx \frac{4}{3}\beta_{Plasma} \approx \beta_{Schock}$$

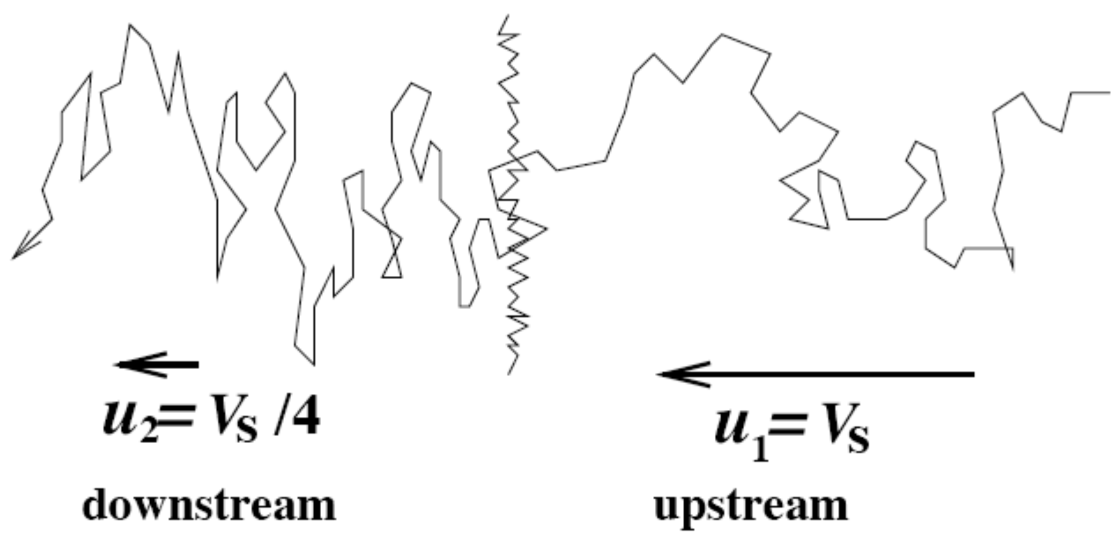
$$\Rightarrow E = E_0 (1+\zeta)^k$$ after k cycles

- Reminder:
  $v_{shock}/v_{material}$ ~ 4/3
  for strong shocks

- The number of cycles depends on the loss rate
  - particles can be lost downstream of the shock

behind the shock the plasma is $v_s/4$ slower than the shock wave itself, particles “diffuse” in the plasma

=> A particle can get lost if it falls to far behind the shock



Energy Spectrum

- Energy gain per cycle

$$\frac{\langle \Delta E \rangle}{E} = \zeta \approx \frac{4}{3}\beta_{Plasma} \approx \beta_{Schock}$$

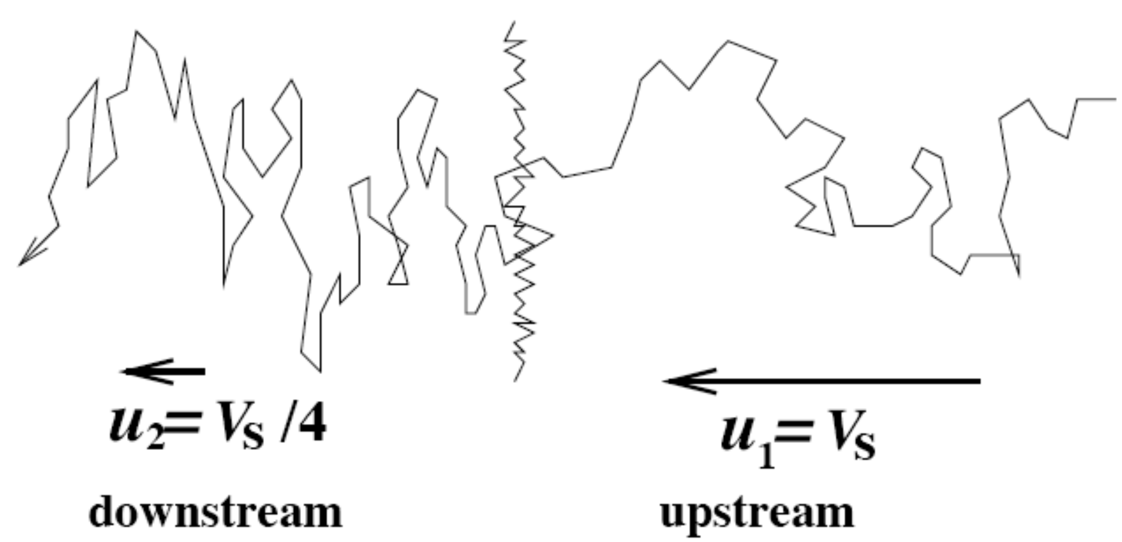
$$\Rightarrow \quad E = E_0 \left(1 + \zeta\right)^k$$ after k cycles

- Reminder:
  $v_{shock}/v_{material}$ ~ 4/3
  for strong shocks

- The number of cycles depends on the loss rate
  - particles can be lost downstream of the shock

behind the shock the plasma is $v_s/4$ slower than the shock wave itself, particles “diffuse” in the plasma

=> A particle can get lost if it falls to far behind the shock



In general: the material behind the shock is 1/4 slower than the front itself, the loss rate depends on that difference:

$$R_{loss} = n_{CR} v_s / 4$$ n_{CR} is the particle density

Energy Spectrum

- Particle flux from "upstream" through the shock:

  Particle movement relative to shock (particle velocity v_t)

$$v_{relativ} = v_s + v_t cos\theta$$

Energy Spectrum

- Particle flux from “upstream” through the shock:

  Particle movement relative to shock (particle velocity v_t)

$$v_{relativ} = v_s + v_t cos\theta$$

  particles cross the shock if $v_{relative} > 0$, which puts constraints on the angle:

$$cos\theta \; > \; -v_s/v_t$$

Energy Spectrum

- Particle flux from “upstream” through the shock:

Particle movement relative to shock (particle velocity v_t)

$$v_{relativ} = v_s + v_t cos\theta$$

particles cross the shock if $v_{relative} > 0$, which puts constraints on the angle:

$$cos\theta \; > \; -v_s/v_t$$

Crossing rate:

n_{CR} is the particle density in the shock region

$$R_{cross} = n_{CR} \frac{1}{4\pi} \int_{-v_s/v_t}^{1} (v_s + v_t cos\theta) 2\pi \, dcos\theta \;\approx\; v_t n_{CR}/4$$

Energy Spectrum

- Particle flux from "upstream" through the shock:

Particle movement relative to shock (particle velocity v_t)

$$v_{relativ} = v_s + v_t cos\theta$$

particles cross the shock if $v_{relative} > 0$, which puts constraints on the angle:

$$cos\theta \; > \; -v_s/v_t$$

Crossing rate:

n_{CR} is the particle density in the shock region

$$R_{cross} = n_{CR}\frac{1}{4\pi}\int_{-v_s/v_t}^{1}(v_s + v_t cos\theta)2\pi \, dcos\theta \;\; \approx \;\; v_t n_{CR}/4$$

Escape probability:

$$P_{escape} \; = \; \frac{R_{loss}}{R_{cross}} \; = \; \frac{v_s}{v_t}$$

Energy Spectrum

- Particle flux from "upstream" through the shock:

Particle movement relative to shock (particle velocity v_t)

$$v_{relativ} = v_s + v_t cos\theta$$

particles cross the shock if $v_{relative} > 0$, which puts constraints on the angle:

$$cos\theta \; > \; -v_s/v_t$$

Crossing rate:

n_{CR} is the particle density in the shock region

$$R_{cross} = n_{CR}\frac{1}{4\pi}\int_{-v_s/v_t}^{1}(v_s + v_t cos\theta)2\pi \, dcos\theta \;\approx\; v_t n_{CR}/4$$

Escape probability:

$$P_{escape} \;=\; \frac{R_{loss}}{R_{cross}} \;=\; \frac{v_s}{v_t}$$

NB: $v_t \sim c$

The probability to cross the shock front at least k times is:

$$P_{cross>k} \;=\; (1 - P_{escape})^k \;=\; \left(1 - \frac{v_s}{v_t}\right)^k \;\approx\; (1 - \beta_{Schock})^k$$

Energy Spectrum

- Reminder: Energy after k cycles:

$$E = E_0\,(1+\zeta)^k \;\approx\; E_0\,(1+\beta_{Schock})^k$$

Energy Spectrum

- Reminder: Energy after k cycles:

$$E = E_0\,(1+\zeta)^k \;\approx\; E_0\,(1+\beta_{Schock})^k$$

$$\Rightarrow\; k \;=\; \frac{ln(E/E_0)}{ln(1+\beta_{Schock})}$$

Energy Spectrum

- Reminder: Energy after k cycles:

$$E = E_0\,(1+\zeta)^k \;\approx\; E_0\,(1+\beta_{Schock})^k$$

$$\Rightarrow\; k \;=\; \frac{ln(E/E_0)}{ln(1+\beta_{Schock})}$$

- Integrated Spectrum (Number of particles with an energy > E):

$$Q(>E) \;\propto\; (1-P_{escape})^k \;=\; (1-\beta_{Schock})^k$$

Energy Spectrum

- Reminder: Energy after k cycles:

$$E = E_0\,(1+\zeta)^k \;\approx\; E_0\,(1+\beta_{Schock})^k$$

$$\Rightarrow\; k \;=\; \frac{ln(E/E_0)}{ln(1+\beta_{Schock})}$$

- Integrated Spectrum (Number of particles with an energy > E):

$$Q(>E) \;\propto\; (1-P_{escape})^k \;=\; (1-\beta_{Schock})^k$$

$$\begin{aligned} lnQ(>E) &= const + k\; ln(1-\beta_{Schock}) \\ &= const' + \frac{ln(1-\beta_{Schock})}{ln(1+\beta_{Schock})}\; lnE \end{aligned}$$

Energy Spectrum

- Reminder: Energy after k cycles:

$$E = E_0\,(1+\zeta)^k \;\approx\; E_0\,(1+\beta_{Schock})^k$$

$$\Rightarrow\; k \;=\; \frac{ln(E/E_0)}{ln(1+\beta_{Schock})}$$

- Integrated Spectrum (Number of particles with an energy > E):

$$Q(>E) \;\propto\; (1-P_{escape})^k \;=\; (1-\beta_{Schock})^k$$

$$\begin{aligned} lnQ(>E) &= const + k\; ln(1-\beta_{Schock}) \\ &= const' + \frac{ln(1-\beta_{Schock})}{ln(1+\beta_{Schock})}\; lnE \end{aligned} \qquad \frac{ln(1-\beta_{Schock})}{ln(1+\beta_{Schock})} \;\approx\; -1$$

Energy Spectrum

- Reminder: Energy after k cycles:

$$E = E_0\,(1+\zeta)^k \;\approx\; E_0\,(1+\beta_{Schock})^k$$

$$\Rightarrow\; k \;=\; \frac{ln(E/E_0)}{ln(1+\beta_{Schock})}$$

- Integrated Spectrum (Number of particles with an energy > E):

$$Q(>E) \;\propto\; (1-P_{escape})^k \;=\; (1-\beta_{Schock})^k$$

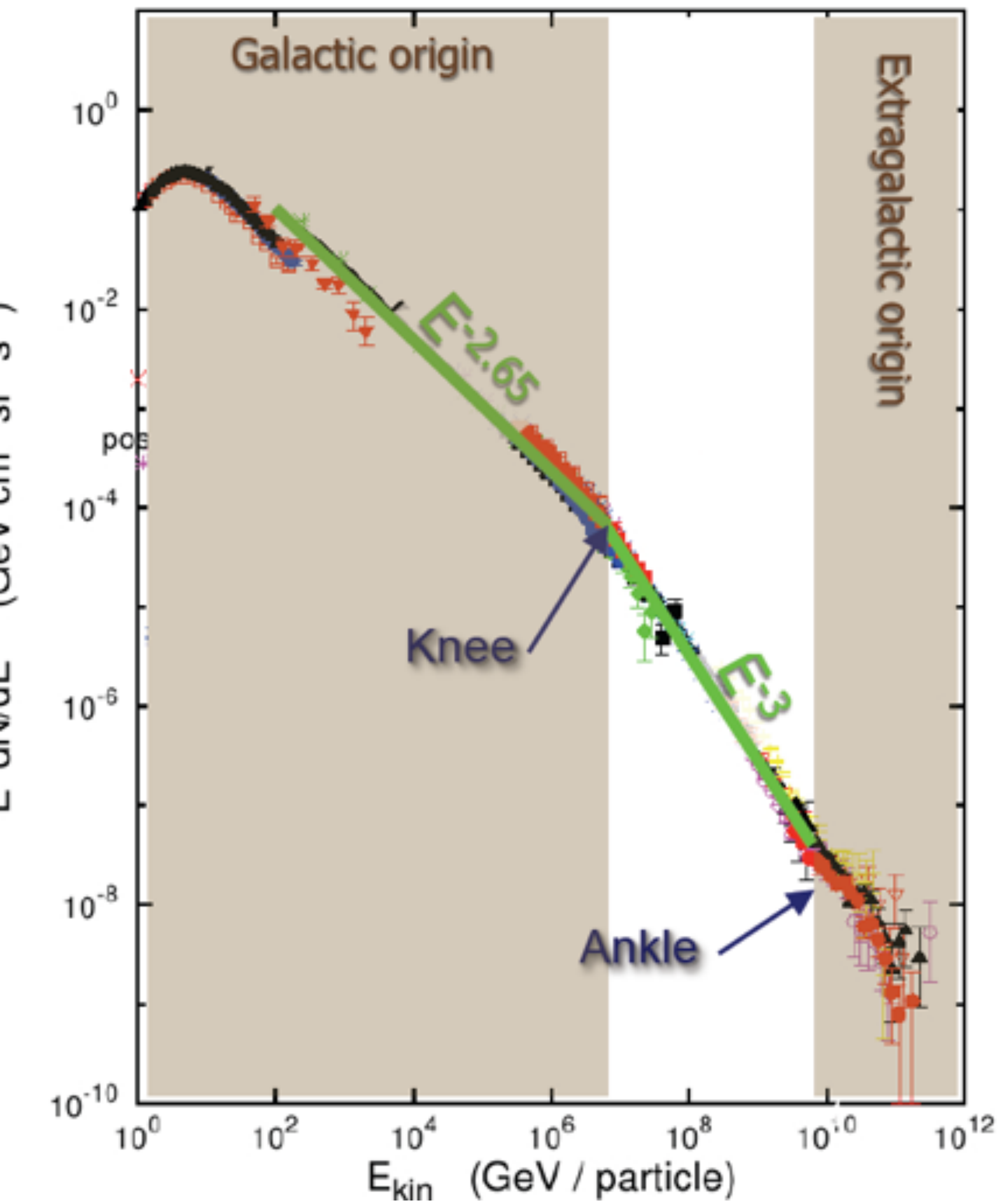
$$\begin{aligned} lnQ(>E) &= const + k\; ln(1-\beta_{Schock}) \\ &= const' + \frac{ln(1-\beta_{Schock})}{ln(1+\beta_{Schock})}\; lnE \end{aligned} \qquad \frac{ln(1-\beta_{Schock})}{ln(1+\beta_{Schock})} \;\approx\; -1$$

- Differential particle spectrum

$$\frac{dN}{dE} \;\propto\; E^{-2}$$

Maximum Energy

- First order Fermi Acceleration can reach energies up to ~ 10^{14} eV for shock waves originating from supernova explosions (incomplete derivation in Backup)



based on a shock lifetime of ~ 1000 years a shock speed of 0.03 c, and a B field in the nT range

➫ Extends up to the knee of the cosmic ray spectrum

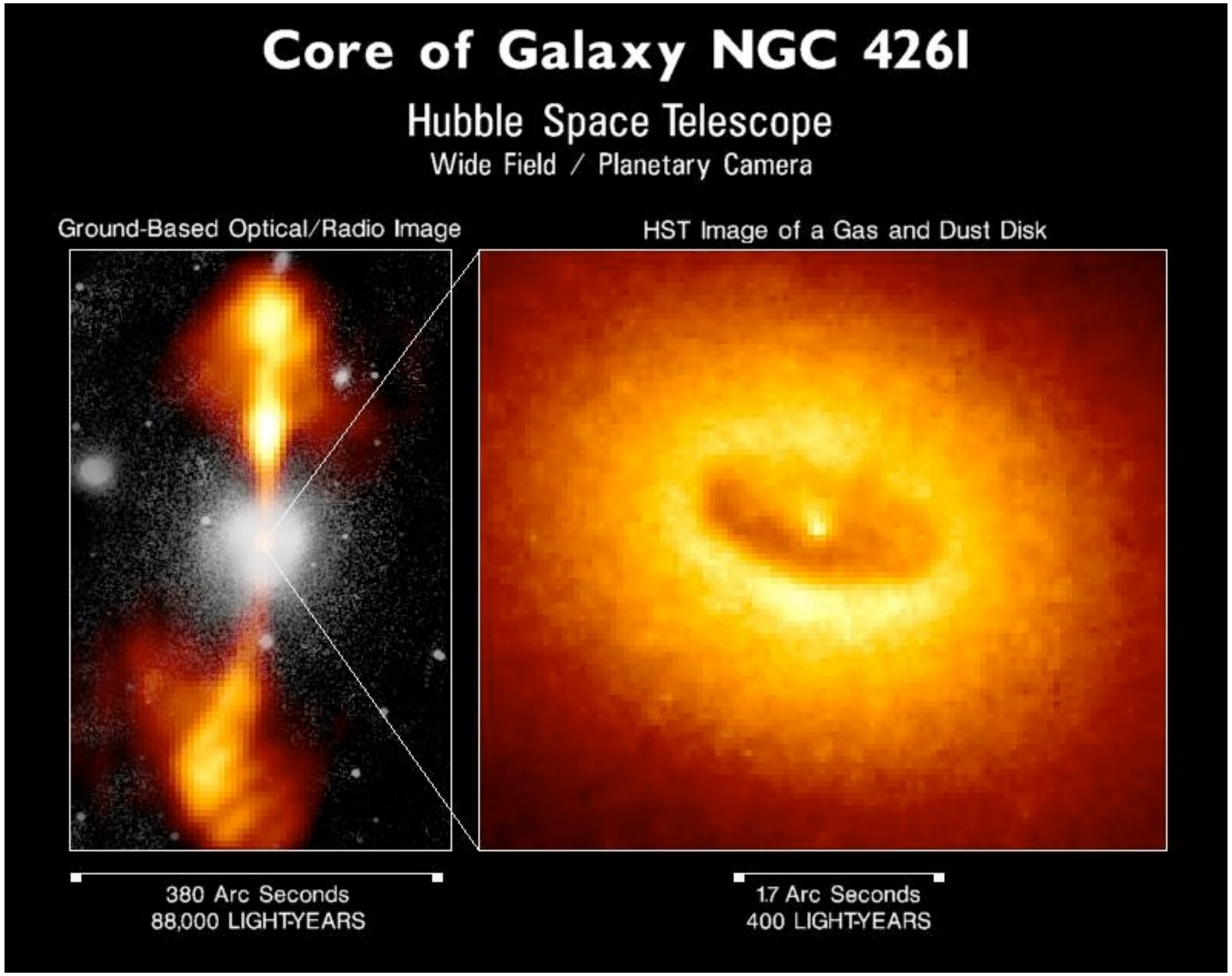
Supernova shock acceleration is well established as a source for cosmic rays

But: What is the origin of the very highest energies above 10^{18} GeV?

S. Coutu, TIPP11

Highest Energies?

- How are energies > 1 PeV reached?

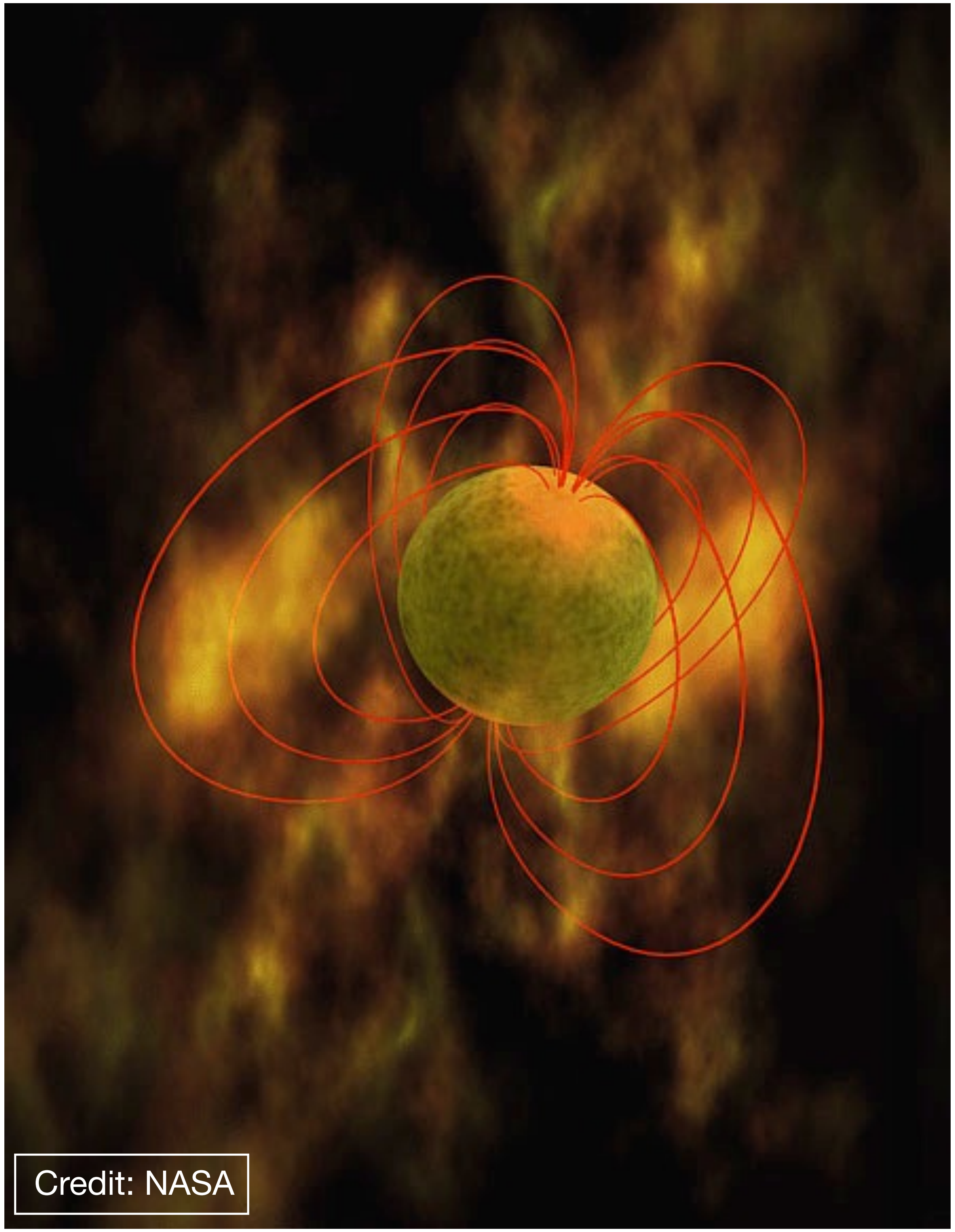


- More energetic events
  - Active galactic nuclei
  - Pulsars (neutron stars)
  - GRB's

➯ Extreme magnetic fields

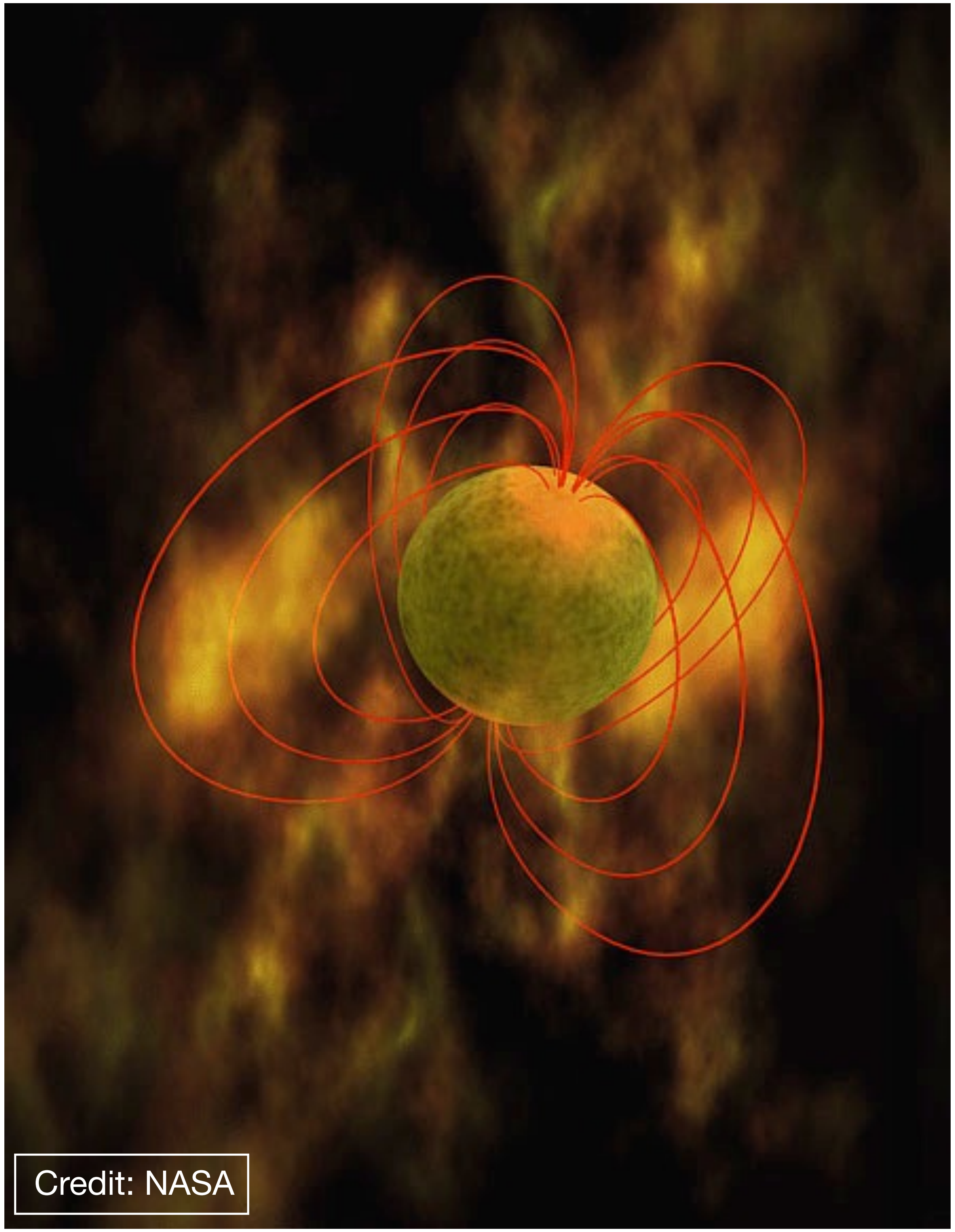
➯ Shock acceleration in highly relativistic jets: additional γ - Factor

One Example: Neutron Stars



- Neutron stars: compact remnants of supernova explosions
  - radius ~ 10 km
  - extreme rotation: up to ~ 40 000 RPM
  - magentic fields up to ~ 10^8 T
  - mass ~ 1.4 M_{Sun}

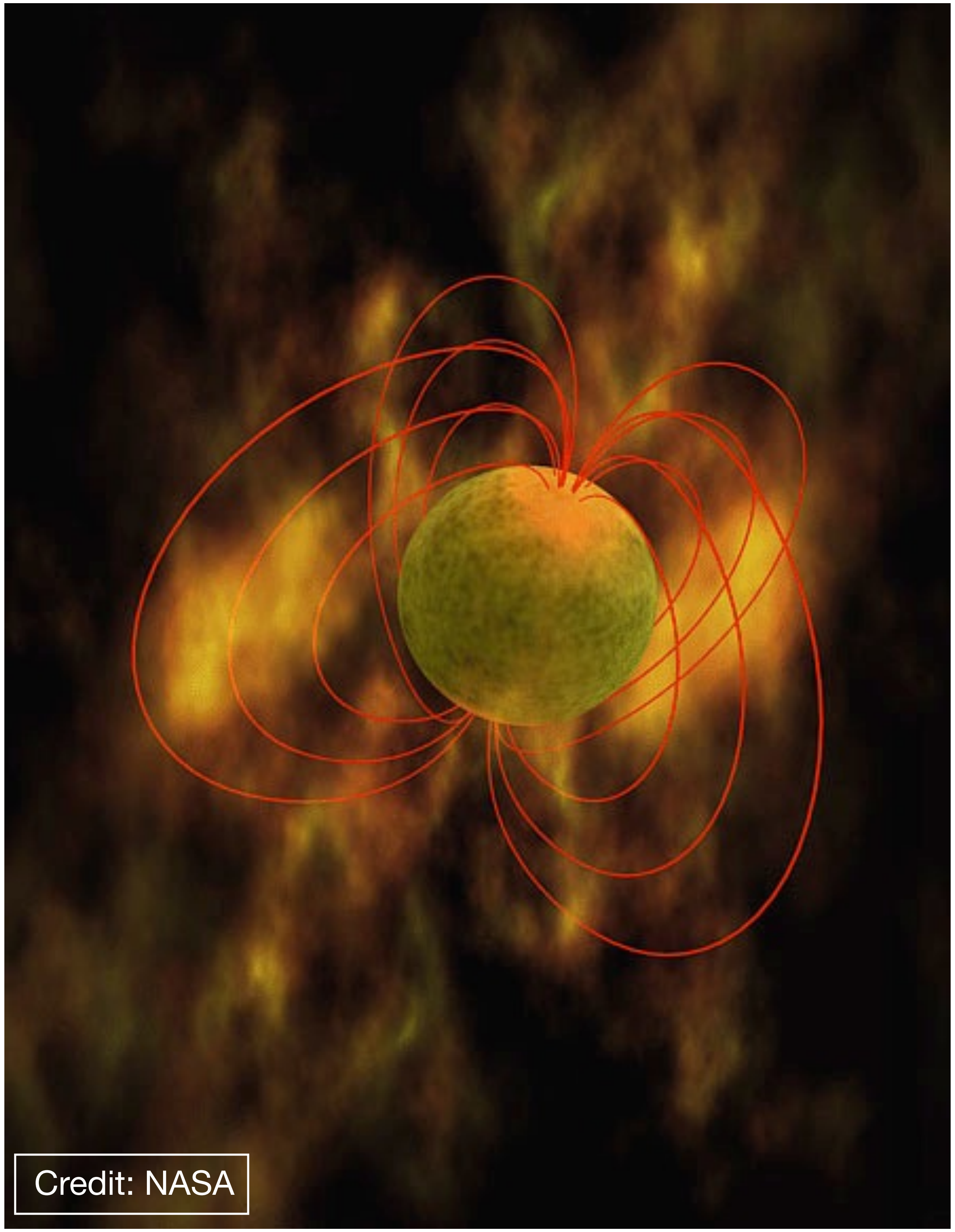
One Example: Neutron Stars



- Neutron stars: compact remnants of supernova explosions
  - radius ~ 10 km
  - extreme rotation: up to ~ 40 000 RPM
  - magentic fields up to ~ 10^8 T
  - mass ~ 1.4 M_{Sun}

Maximum energy: limited by the Larmor radius of the particle: The particle will escape if r_L gets to large!

One Example: Neutron Stars

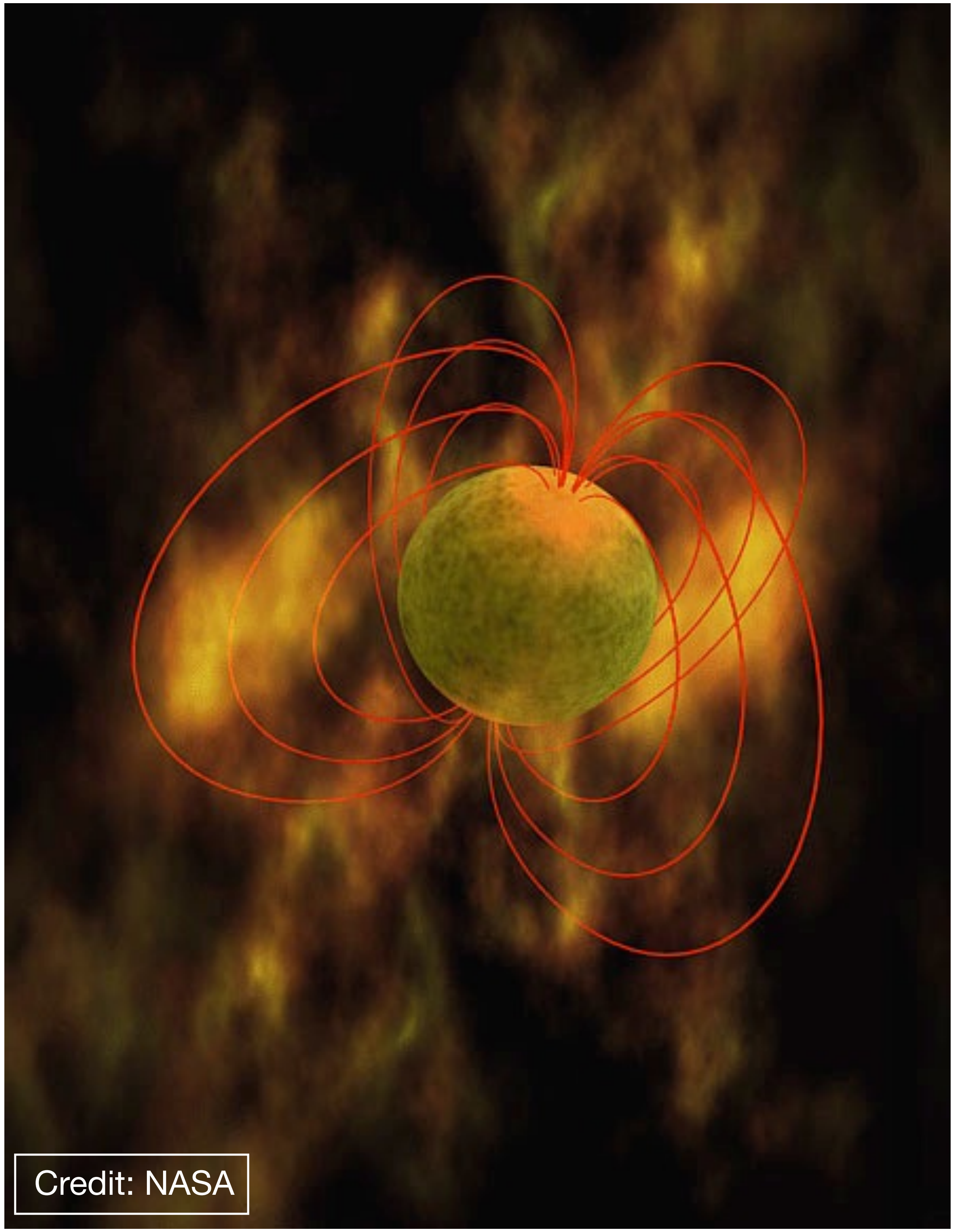


- Neutron stars: compact remnants of supernova explosions
  - radius ~ 10 km
  - extreme rotation: up to ~ 40 000 RPM
  - magentic fields up to ~ 10^8 T
  - mass ~ 1.4 M_{Sun}

Maximum energy: limited by the Larmor radius of the particle: The particle will escape if r_L gets to large!

$$r_L = \frac{p_T}{|q|B} \qquad p_T \approx E/c$$

One Example: Neutron Stars



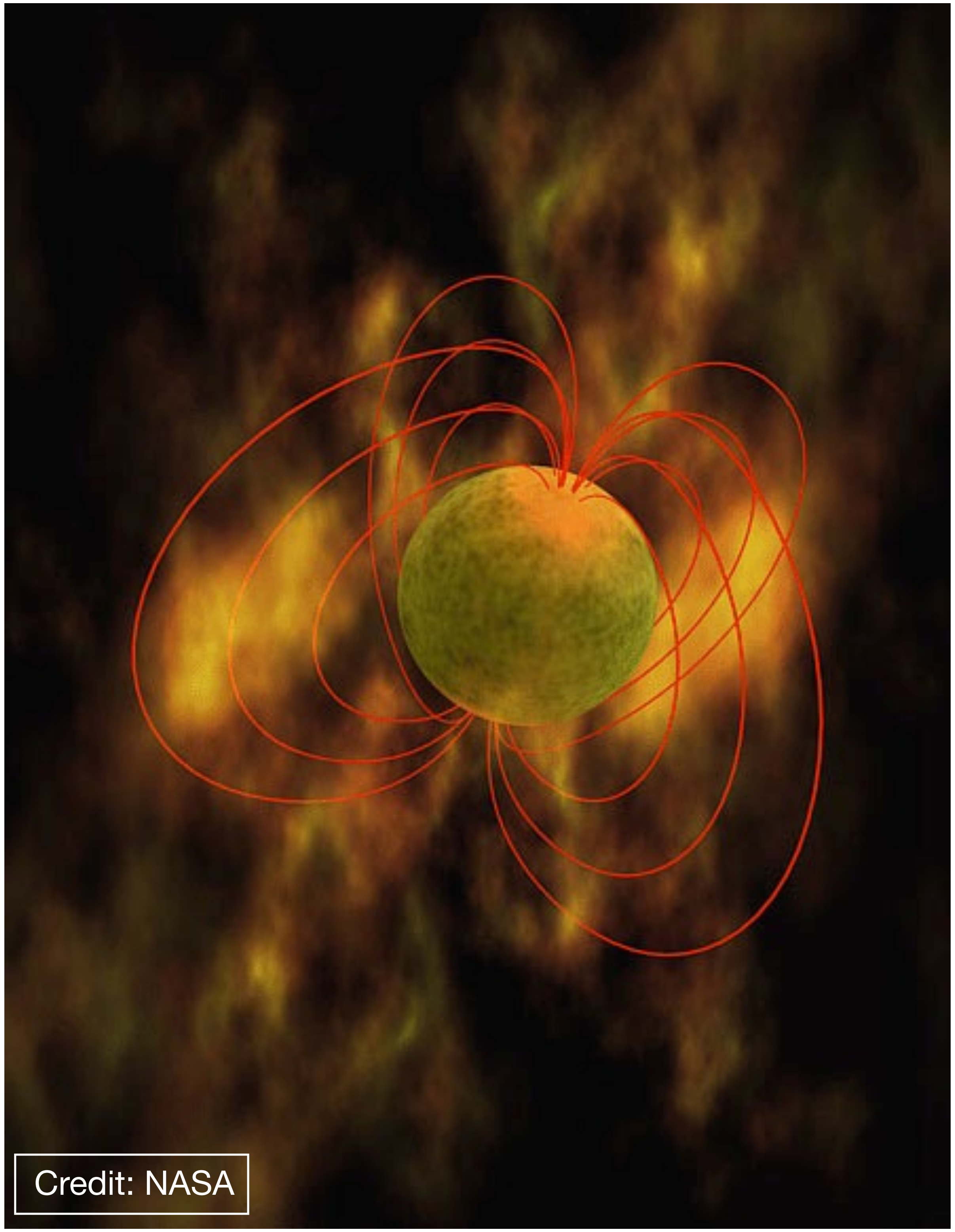
- Neutron stars: compact remnants of supernova explosions
  - radius ~ 10 km
  - extreme rotation: up to ~ 40 000 RPM
  - magentic fields up to ~ 10^8 T
  - mass ~ 1.4 M_{Sun}

Maximum energy: limited by the Larmor radius of the particle: The particle will escape if r_L gets to large!

$$r_L = \frac{p_T}{|q|B} \qquad p_T \approx E/c$$

➩ $$E_{max} = Z\, e\, c\, B\, R$$

One Example: Neutron Stars



- Neutron stars: compact remnants of supernova explosions
  - radius ~ 10 km
  - extreme rotation: up to ~ 40 000 RPM
  - magentic fields up to ~ 10^8 T
  - mass ~ 1.4 M_{Sun}

Maximum energy: limited by the Larmor radius of the particle: The particle will escape if r_L gets to large!

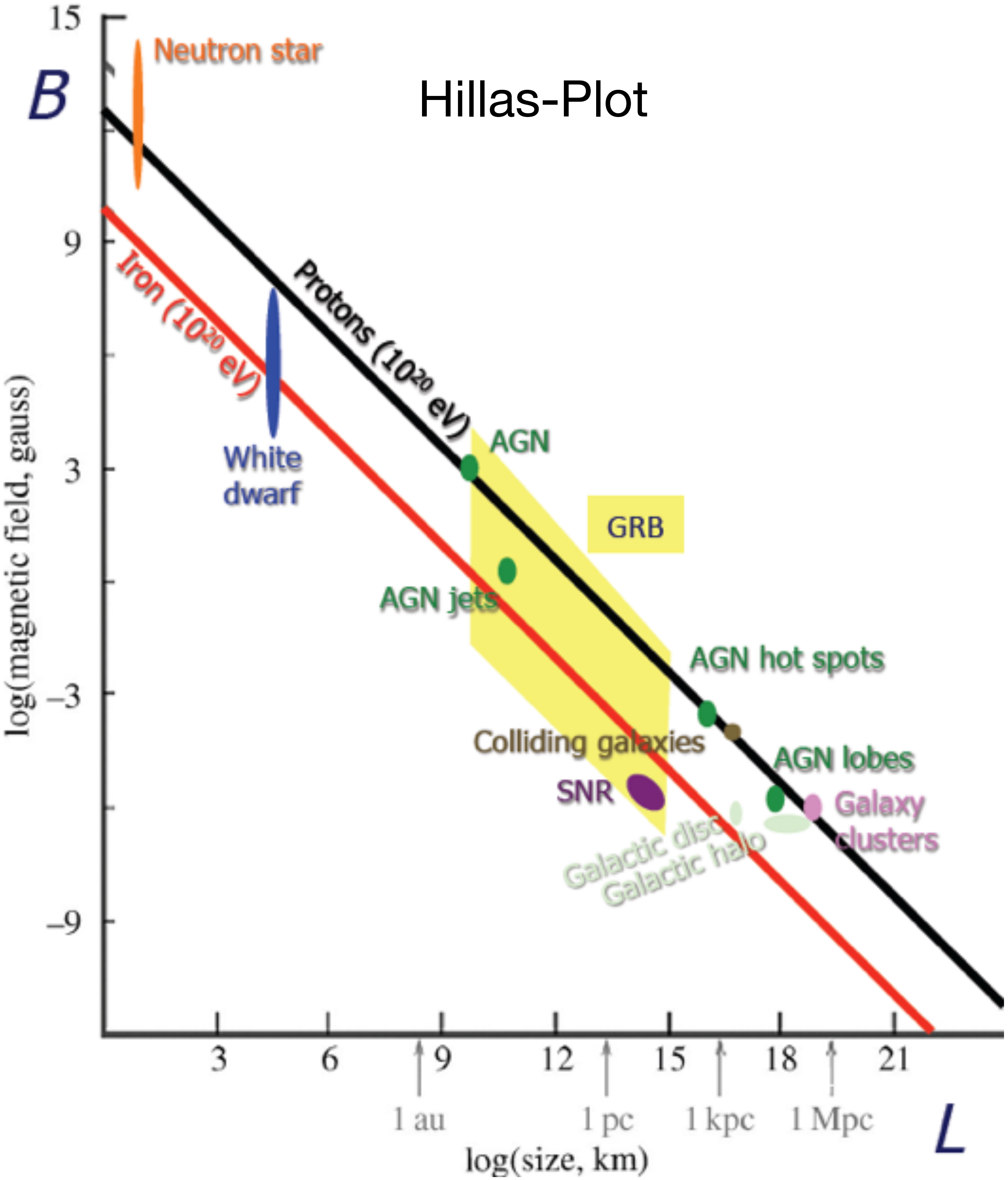
$$r_L = \frac{p_T}{|q|B} \qquad p_T \approx E/c$$

➪ $$E_{max} = Z\, e\, c\, B\, R$$

For Z = 1 (protons):

E_{max} ~ 5 J = 3 x 10^{20} eV

Candidates for the Highest Energies



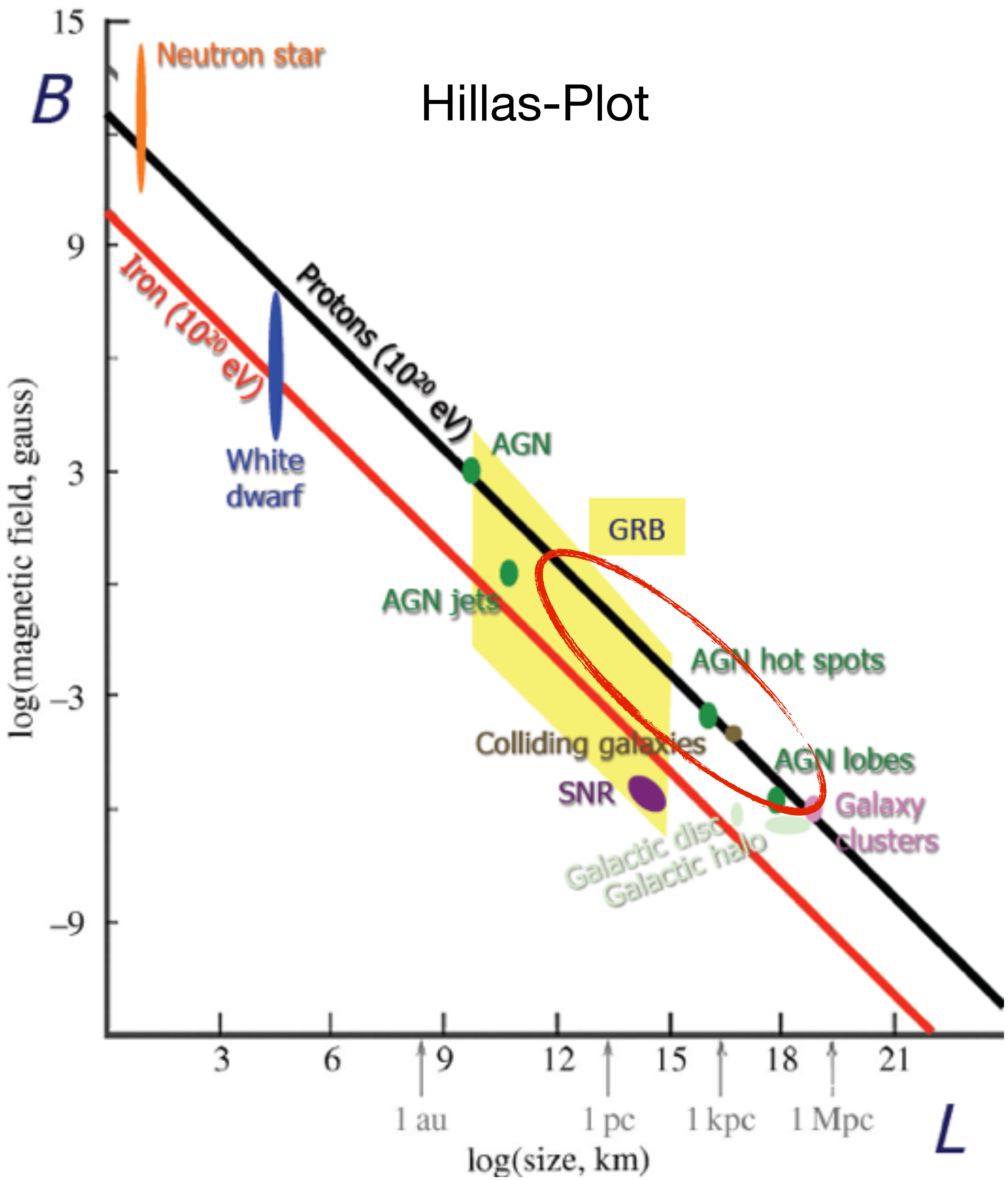
Particles are accelerated as long as they stay in the high field region: $r_L < L$

$$E_{max} \approx 10^{20} eV\, Z B_{\mu G} L_{100\,kpc}$$

For Iron nuclei 26 x higher energies are possible relative to protons!

S. Coutu, TIPP11

Candidates for the Highest Energies



Particles are accelerated as long as they stay in the high field region: $r_L < L$

$$E_{max} \approx 10^{20} eV\ Z B_{\mu G} L_{100 kpc}$$

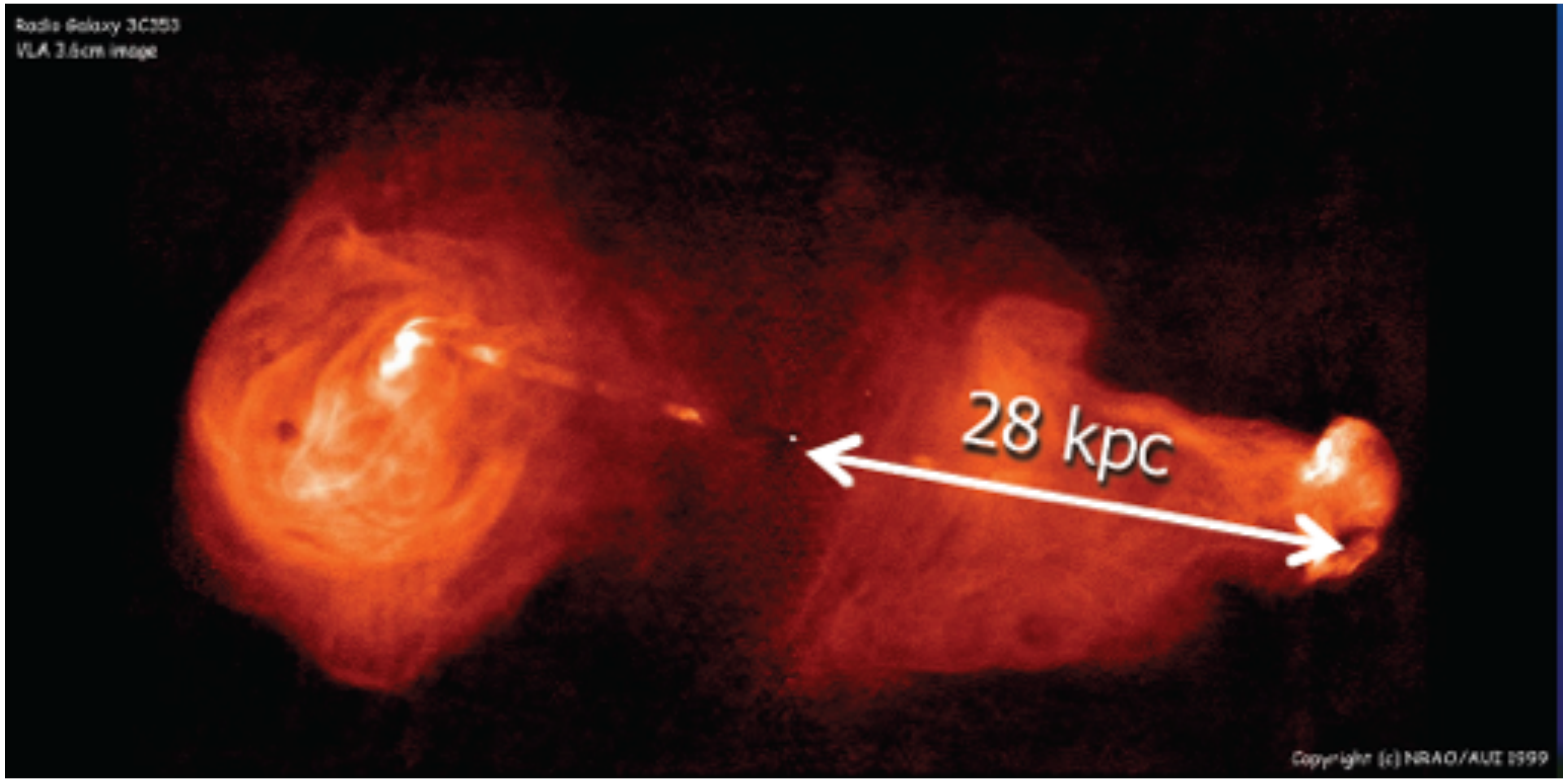
For Iron nuclei 26 x higher energies are possible relative to protons!

Beyond this simple consideration: Radiation losses in the source have to be taken into account: synchotron radiation, photo reactions

S. Coutu, TIPP11

One Example

- 3C353 - Active galaxy 130 Mpc away



… more about the highest later in the lecture!

Propagation of Cosmic Rays

- The source spectrum of shock acceleration follows an E^{-2} distribution, but we observe $E^{-2.7}$, why?

▸ Energy-dependent loss of particles when travelling through the galaxy

- Important contributions:
  - diffusion
  - convection
  - acceleration

transport in turbulent galactic magnetic fields

  - decay of unstable particles and nuclei
  - collisions
  - cascade production, spallation of heavy nuclei

loss processes

Leaky Box Model

- Very simple model assuming cosmic rays propagate freely in the galaxy with a constant escape / loss probability

$$N(E,t) = N_0(E)e^{-t/\tau_{escape}}$$

Leaky Box Model

- Very simple model assuming cosmic rays propagate freely in the galaxy with a constant escape / loss probability

$$N(E,t) = N_0(E)e^{-t/\tau_{escape}}$$

- Adaptation to reality - escape / loss probability is energy dependent:
  - particles with higher energy can more easily leave the “magnetic confinement” in the galaxy and are more likely to participate in inelastic reactions
  - ‣ The observed spectrum gets steeper than the source spectrum: $E^{-2.7}$

Leaky Box Model

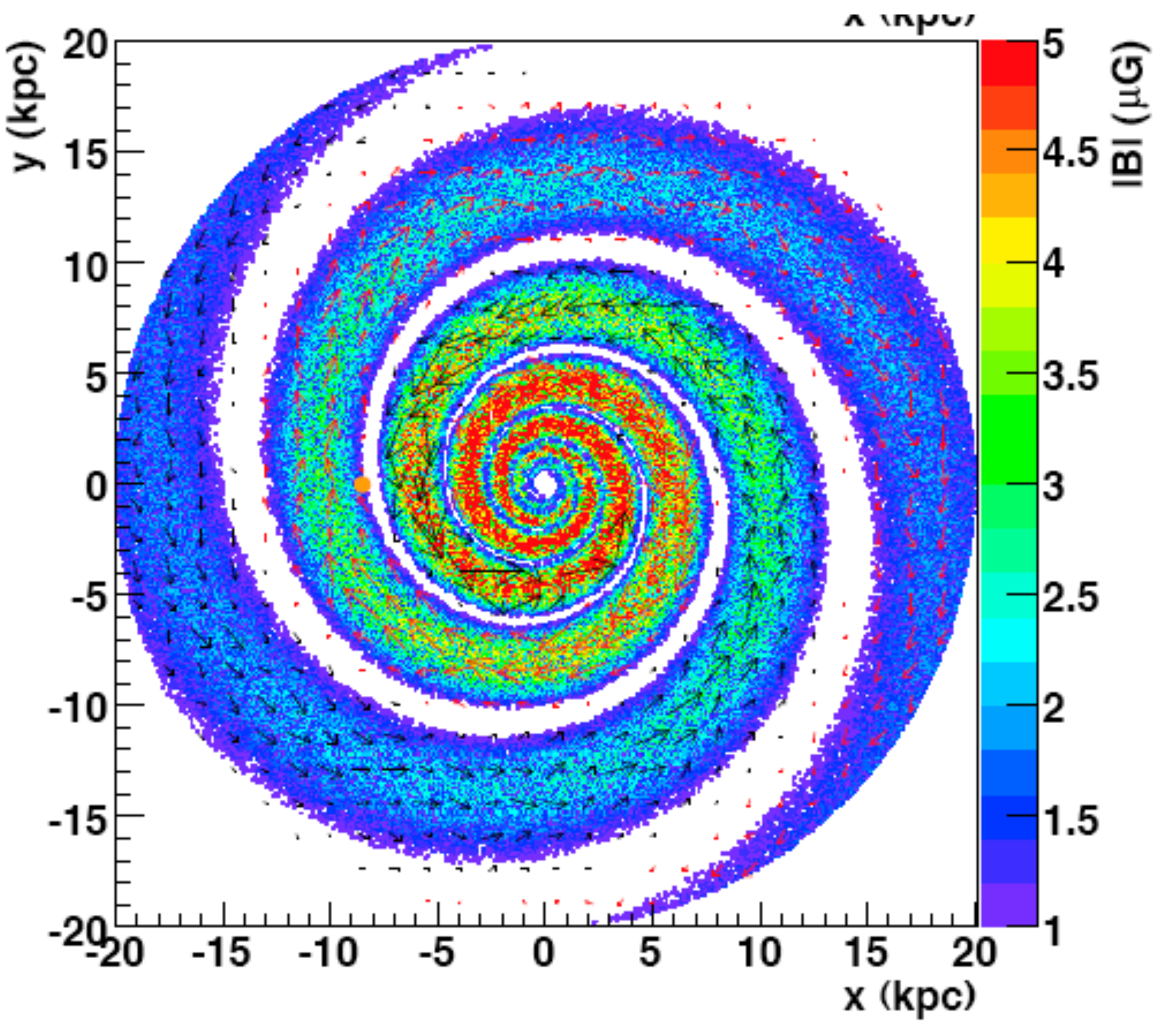
- Very simple model assuming cosmic rays propagate freely in the galaxy with a constant escape / loss probability

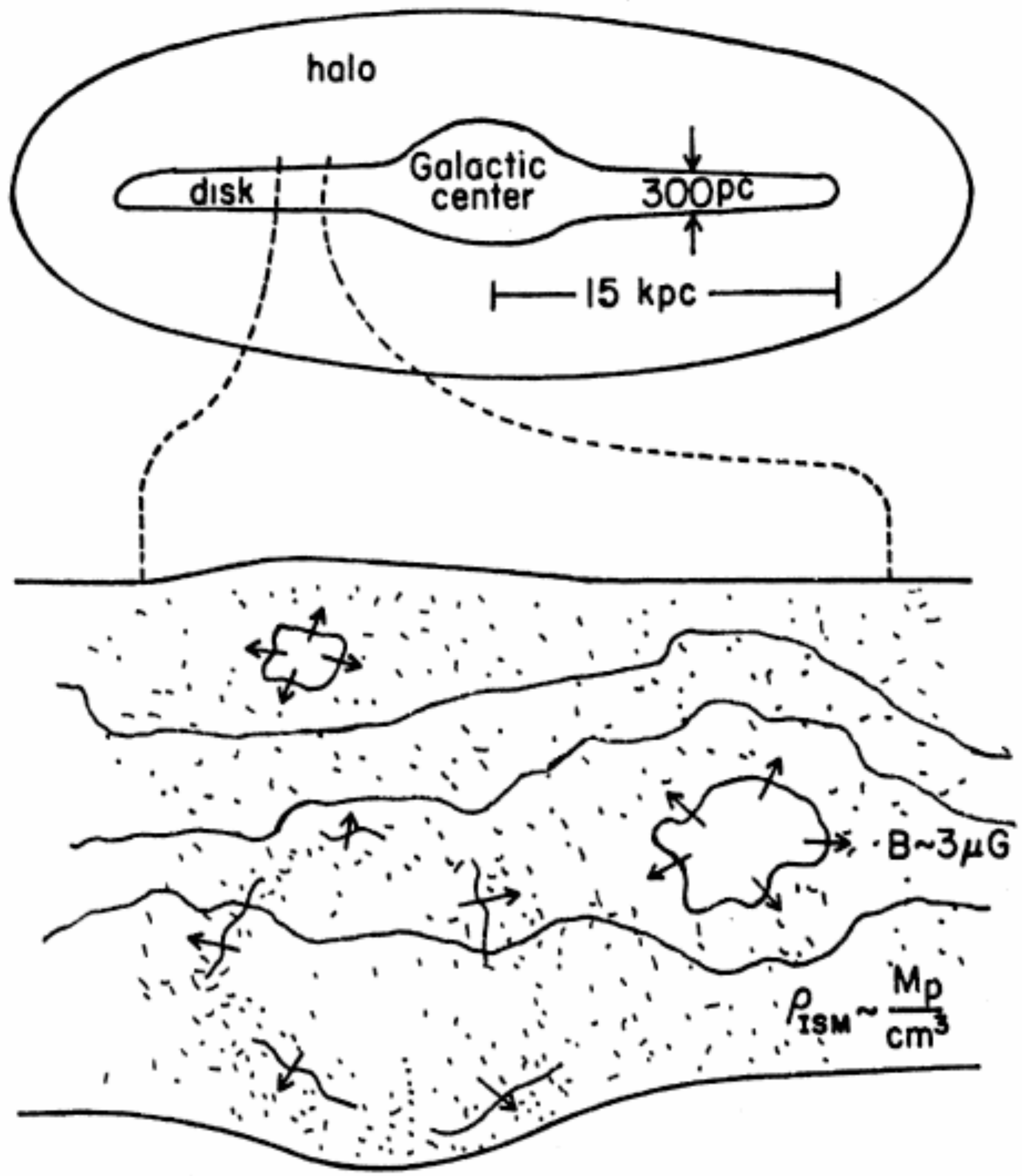
$$N(E,t) = N_0(E)e^{-t/\tau_{escape}}$$

- Adaptation to reality - escape / loss probability is energy dependent:
  - particles with higher energy can more easily leave the “magnetic confinement” in the galaxy and are more likely to participate in inelastic reactions
  - ▸ The observed spectrum gets steeper than the source spectrum: $E^{-2.7}$
- Loss probability due to inelastic reactions depends on amount of traversed matter
  - Density of the ISM in the galaxy: ~ 1 proton/cm^3 ~ 1.7 x 10^{-24} g/cm^3
  - ▸ per year one particle traverses ~ 1.5 x10^{-6} g/cm^2
  - ▸ loss after traversing ~ 5 - 10 g/cm^2 (derived from observed composition)
  - ▸ Particles stay in the galaxy for about 5 x 10^6 years

Magnetic Fields in the Galaxy

- Magnetic field in the galaxy along spiral arms, with additional turbulent contributions overlaid
- typical strength ~0.1 nT

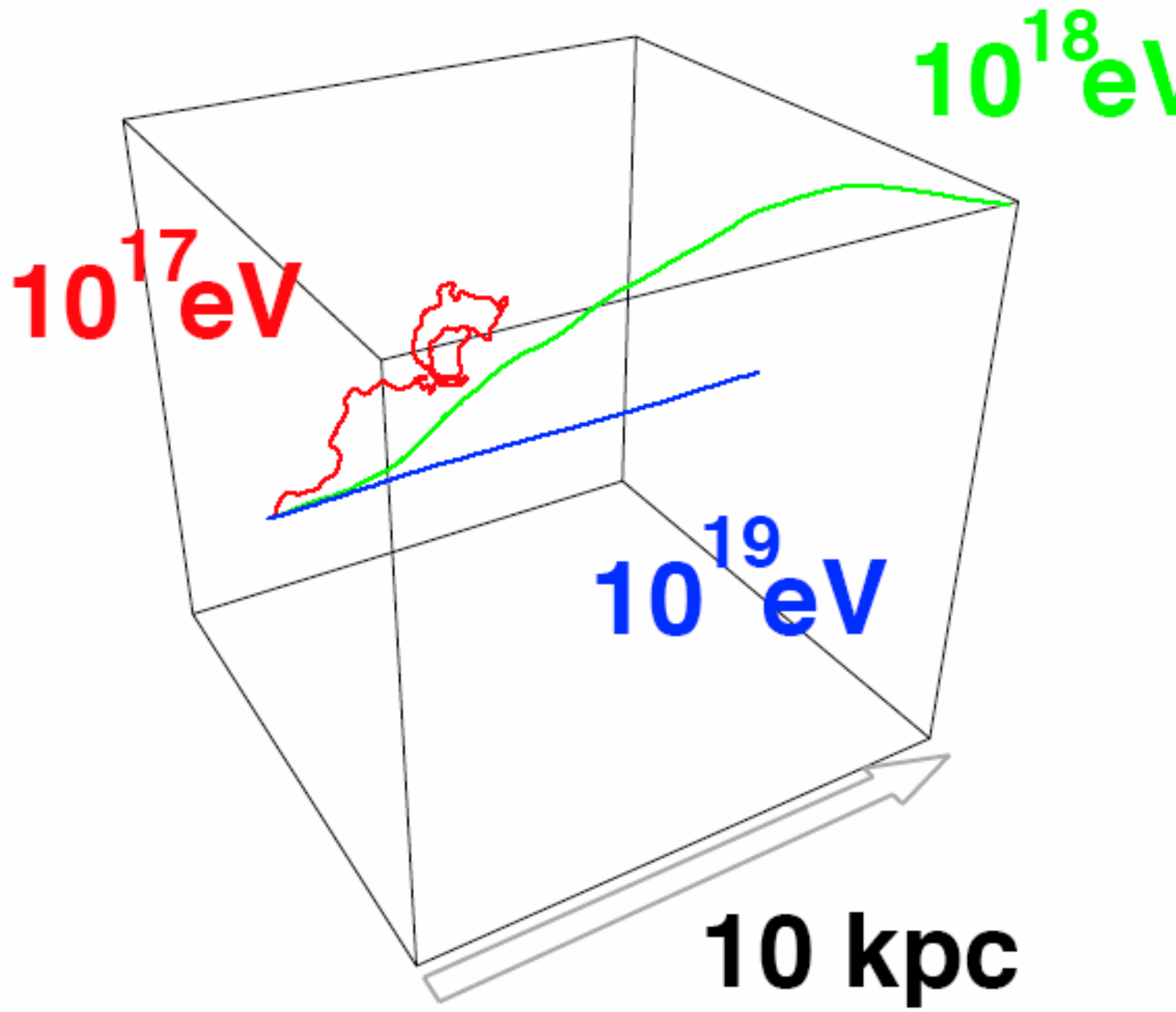




Propagation of Particles in Magnetic Fields

- Charged particles are deflected by cosmic magnetic fields

- To demonstrate: toy simulation with magnetic fields of ~0.1 nT and a coherence length of ~100 pc
    - Particles start from the left center with different energies



- Only the very highest energies (E ~ 10^{19} eV) can show the way to their sources - all other particles get substantially deflected and arrive from random directions

Summary

- Cosmic rays are known since 100 years
  - Discovered by Victor Hess on balloon flights

- Acceleration mechanism via scattering on randomly moving cosmic clouds proposed by Fermi 60 years ago (second order Fermi Acceleration)
  - A proof of principle, but insufficient to reach high energies

- Acceleration in shock fronts created by supernovae (first order Fermi Acceleration) can explain energies at least up to ~10^{14} eV
  - Shock front acceleration is most likely also responsible for even higher energies in objects such as pulsars, blazars, AGNs...

Summary

- Cosmic rays are known since 100 years
  - Discovered by Victor Hess on balloon flights

- Acceleration mechanism via scattering on randomly moving cosmic clouds proposed by Fermi 60 years ago (second order Fermi Acceleration)
  - A proof of principle, but insufficient to reach high energies

- Acceleration in shock fronts created by supernovae (first order Fermi Acceleration) can explain energies at least up to ~10^{14} eV
  - Shock front acceleration is most likely also responsible for even higher energies in objects such as pulsars, blazars, AGNs...

Next Lecture: 15.05., “Detectors in Particle & Astroparticle Physics”, F. Simon

Lecture Overview

24.04.	Introduction & Accelerators
01.05.	**Holiday - No Lecture**
08.05	Cosmic Accelerators
15.05.	Detectors
22.05.	The Standard Model
29.05.	QCD and Jets
05.06.	**Holiday - No Lecture**
12.06.	Neutrinos I
19.06.	Neutrinos II
26.06	**most likely: No Lecture**
03.07.	Cosmic Rays I
10.07.	Cosmic Rays II
17.07.	Precision Experiments
24.07.	Dark Matter, Dark Energy & Gravitational Waves

Backup

Erreichbare Energie

- Die Rate des Energiezuwaches ist gegeben durch die Dauer eines Zyklus und durch den Zuwachs pro Zyklus:

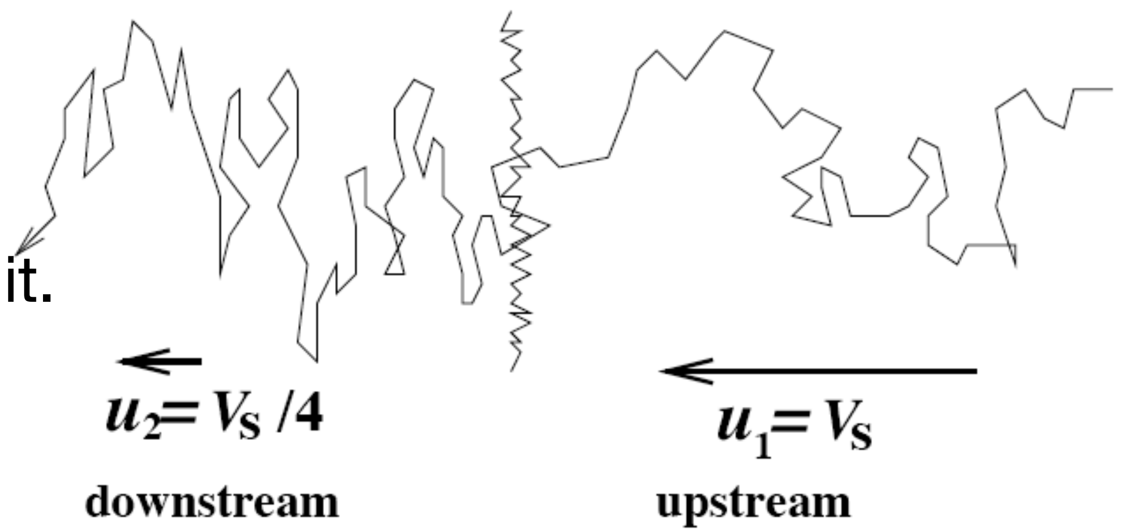
$$\frac{dE}{dt} = \frac{\Delta E}{t_{cycle}} = \frac{E\,\beta_{Schock}}{t_{cycle}}$$

Erreichbare Energie

- Die Rate des Energiezuwaches ist gegeben durch die Dauer eines Zyklus und durch den Zuwachs pro Zyklus:

$$\frac{dE}{dt} = \frac{\Delta E}{t_{cycle}} = \frac{E\,\beta_{Schock}}{t_{cycle}}$$

- Betrachtung im Bezugssystem des Schocks:



Hinter dem Schock:
Teilchen diffundiert (Diffusionskoeffizient k_2),
und “fliesst” mit der Plasmageschwindigkeit mit.
Verweildauer hinter dem Schock: t

Diffusion: $\sqrt{k_2 t}$

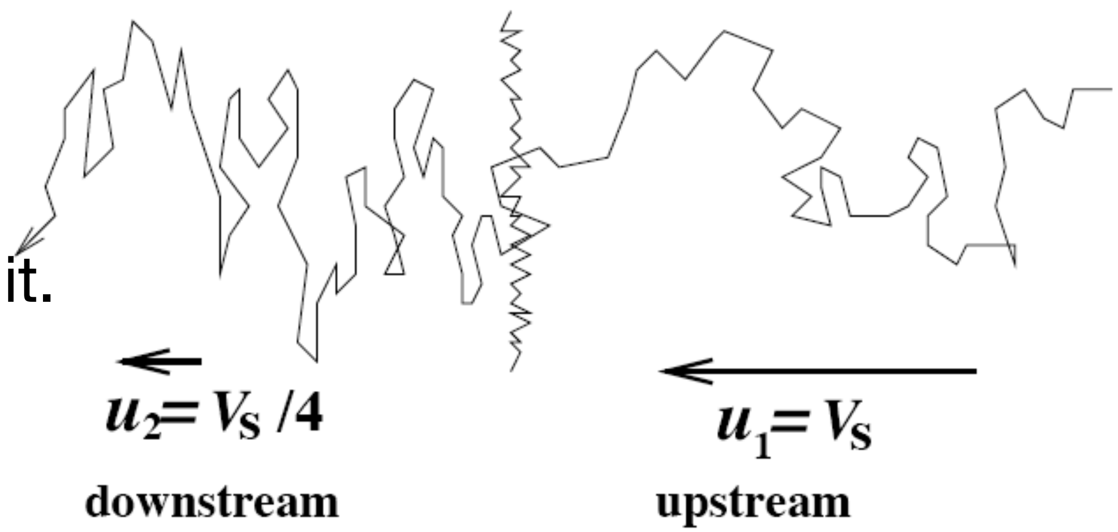
gerichtete Bewegung: $u_2 t$

Erreichbare Energie

- Die Rate des Energiezuwaches ist gegeben durch die Dauer eines Zyklus und durch den Zuwachs pro Zyklus:

$$\frac{dE}{dt} = \frac{\Delta E}{t_{cycle}} = \frac{E\,\beta_{Schock}}{t_{cycle}}$$

- Betrachtung im Bezugssystem des Schocks:



Hinter dem Schock:
Teilchen diffundiert (Diffusionskoeffizient k_2),
und “fliesst” mit der Plasmageschwindigkeit mit.
Verweildauer hinter dem Schock: t

Diffusion: $\sqrt{k_2 t}$

gerichtete Bewegung: $u_2 t$

Hohe Wahrscheinlichkeit, wieder in den Schock zu kommen: $\sqrt{k_2 t} \gg u_2 t$

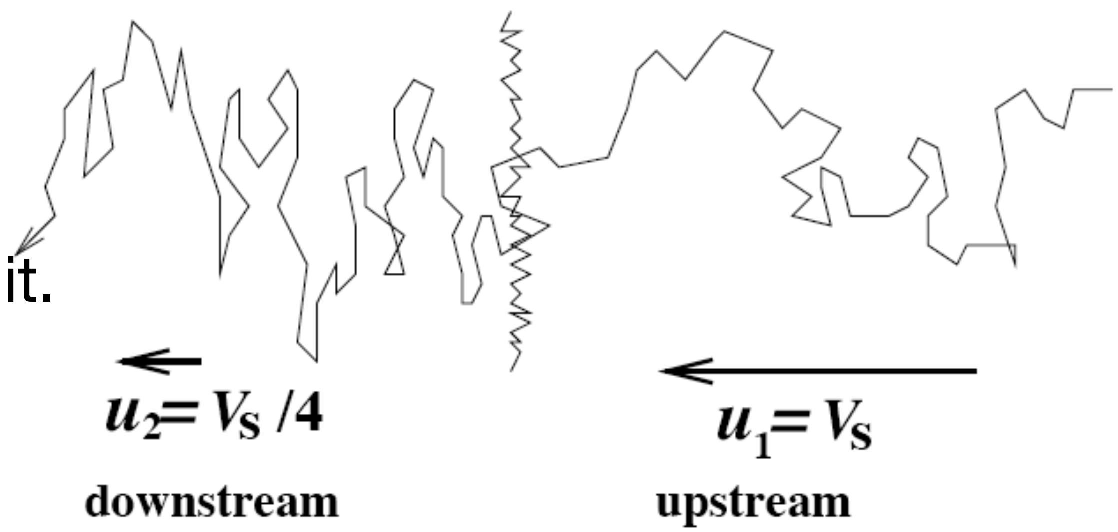
Erreichbare Energie

- Die Rate des Energiezuwaches ist gegeben durch die Dauer eines Zyklus und durch den Zuwachs pro Zyklus:

$$\frac{dE}{dt} = \frac{\Delta E}{t_{cycle}} = \frac{E\,\beta_{Schock}}{t_{cycle}}$$

- Betrachtung im Bezugssystem des Schocks:

Hinter dem Schock:
Teilchen diffundiert (Diffusionskoeffizient k_2),
und “fliesst” mit der Plasmageschwindigkeit mit.
Verweildauer hinter dem Schock: t



Diffusion: $\sqrt{k_2 t}$

gerichtete Bewegung: $u_2 t$

Hohe Wahrscheinlichkeit, wieder in den Schock zu kommen: $\sqrt{k_2 t} \gg u_2 t$

Hohe Wahrscheinlichkeit, den Schock für immer zu verlassen: $\sqrt{k_2 t} \ll u_2 t$

Erreichbare Energie

- "Grenze", die entscheidet, ob ein Teilchen "verloren" ist: $\sim \; k_2/u_2$

- Verweildauer hinter dem Schock aus Teilchendichte und Übergangsrate:

$$t_2 \approx \; n_{CR} \frac{k_2}{u_2} \frac{1}{R_{Cross}} = \frac{4}{c} \frac{k_2}{u_2}$$

Erreichbare Energie

- “Grenze”, die entscheidet, ob ein Teilchen “verloren” ist: $\sim \; k_2/u_2$

- Verweildauer hinter dem Schock aus Teilchendichte und Übergangsrate:

$$t_2 \approx \; n_{CR} \frac{k_2}{u_2} \frac{1}{R_{Cross}} = \frac{4}{c} \frac{k_2}{u_2}$$

- Analoge Überlegungen für die “Upstream” - Zone:
  - k_1/u_1 markiert hier die Grenze zwischen Teilchen, die schon von hinter dem Schock gekommen und und solchen, die noch nie durch den Schock gegangen sind

$$t_1 \approx \; \frac{4}{c} \frac{k_1}{u_1}$$

Erreichbare Energie

- "Grenze", die entscheidet, ob ein Teilchen "verloren" ist: $\sim \; k_2/u_2$

- Verweildauer hinter dem Schock aus Teilchendichte und Übergangsrate:

$$t_2 \approx \; n_{CR} \frac{k_2}{u_2} \frac{1}{R_{Cross}} = \frac{4}{c} \frac{k_2}{u_2}$$

- Analoge Überlegungen für die "Upstream" - Zone:
    - k_1/u_1 markiert hier die Grenze zwischen Teilchen, die schon von hinter dem Schock gekommen und und solchen, die noch nie durch den Schock gegangen sind

$$t_1 \approx \; \frac{4}{c} \frac{k_1}{u_1}$$

- Damit ergibt sich die Zyklus-Dauer:

$$t_{cycle} \; = \; t_1 + t_2 \; \approx \; \frac{4}{c} \left(\frac{k_1}{u_1} + \frac{k_2}{u_2} \right) = \frac{4}{\beta_{Schock} c^2} (k_1 + 4k_2)$$

Erreichbare Energie

- Die Diffusionskonstante hängt vom Magnetfeld ab (Diffusionslänge muss mindestens so groß sein wie der Lamor-Radius, damit Streuung an Magnetfeldänderungen funktioniert) - Bohm-Diffusions-Koeffizient:

$$k = \frac{1}{3} r_L \, c \ , \ r_L = \frac{p}{ZeB} \sim \frac{E}{cZeB}$$

Erreichbare Energie

- Die Diffusionskonstante hängt vom Magnetfeld ab (Diffusionslänge muss mindestens so groß sein wie der Lamor-Radius, damit Streuung an Magnetfeldänderungen funktioniert) - Bohm-Diffusions-Koeffizient:

$$k = \frac{1}{3} r_L \, c \; , \; r_L = \frac{p}{ZeB} \sim \frac{E}{cZeB}$$

- Es folgt für die Zyklus-Dauer für $k_1 = k_2$

$$t_{cycle} = \frac{20}{3} \frac{E}{c^2 \beta_{Schock} ZeB} \propto E$$

Erreichbare Energie

- Die Diffusionskonstante hängt vom Magnetfeld ab (Diffusionslänge muss mindestens so groß sein wie der Lamor-Radius, damit Streuung an Magnetfeldänderungen funktioniert) - Bohm-Diffusions-Koeffizient:

$$k = \frac{1}{3} r_L \, c \ , \ r_L \ = \ \frac{p}{ZeB} \sim \frac{E}{cZeB}$$

- Es folgt für die Zyklus-Dauer für $k_1 = k_2$

$$t_{cycle} \ = \ \frac{20}{3} \frac{E}{c^2 \beta_{Schock} ZeB} \ \propto \ E$$

- Erreichbare Energie:

$$E_{max} \ = \ \int_0^{t_{acc}} \frac{dE}{dt} dt = \int_0^{t_{acc}} \frac{E\beta_{Schock}}{\frac{20E}{3c^2 \beta_{Schock} ZeB}} = \frac{3}{20} \beta^2_{Schock} c^2 ZeBt_{acc}$$

Erreichbare Energie

- Die Diffusionskonstante hängt vom Magnetfeld ab (Diffusionslänge muss mindestens so groß sein wie der Lamor-Radius, damit Streuung an Magnetfeldänderungen funktioniert) - Bohm-Diffusions-Koeffizient:

$$k = \frac{1}{3} r_L \, c \;, \; r_L = \frac{p}{ZeB} \sim \frac{E}{cZeB}$$

- Es folgt für die Zyklus-Dauer für $k_1 = k_2$

$$t_{cycle} = \frac{20}{3} \frac{E}{c^2 \beta_{Schock} ZeB} \propto E$$

- Erreichbare Energie:

$$E_{max} = \int_0^{t_{acc}} \frac{dE}{dt} dt = \int_0^{t_{acc}} \frac{E \beta_{Schock}}{\frac{20E}{3c^2 \beta_{Schock} ZeB}} = \frac{3}{20} \beta_{Schock}^2 c^2 ZeB t_{acc}$$

- Für typische Werte (β_{Schock} ~ 0.03, B ~ 0.3 nT, t_{acc} ~ 1000 Jahre) E_{max} ~ 10^{14} eV (für Protonen)

‣ bis zum Knie der Verteilung