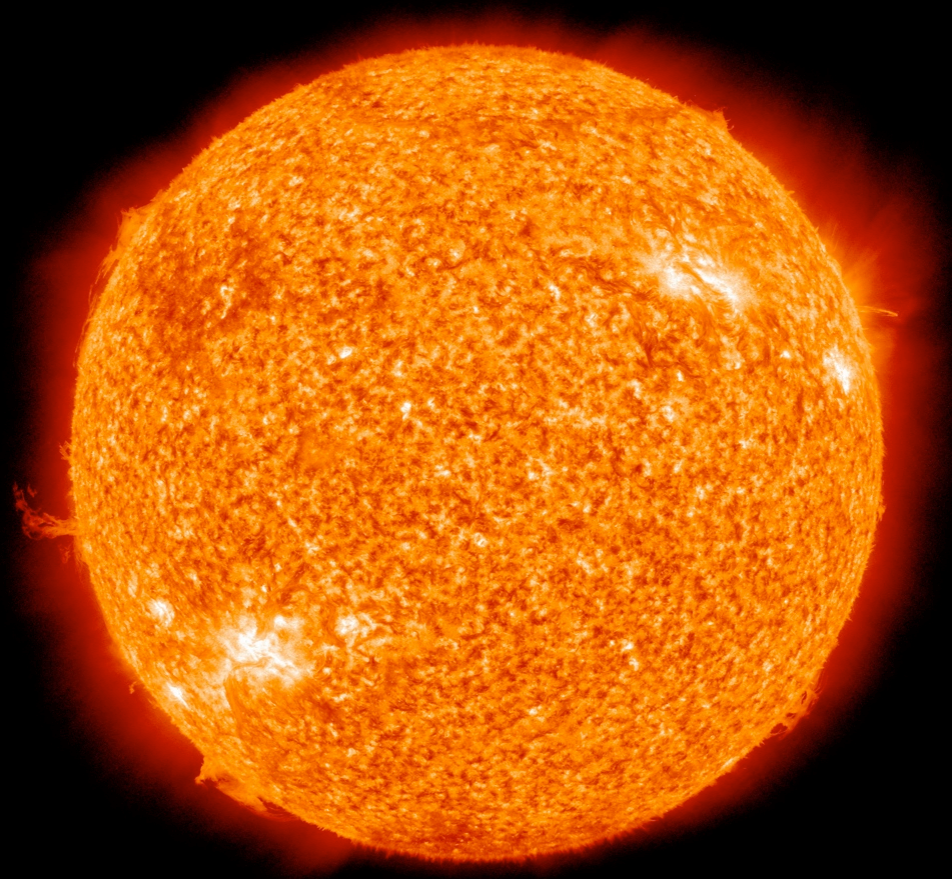


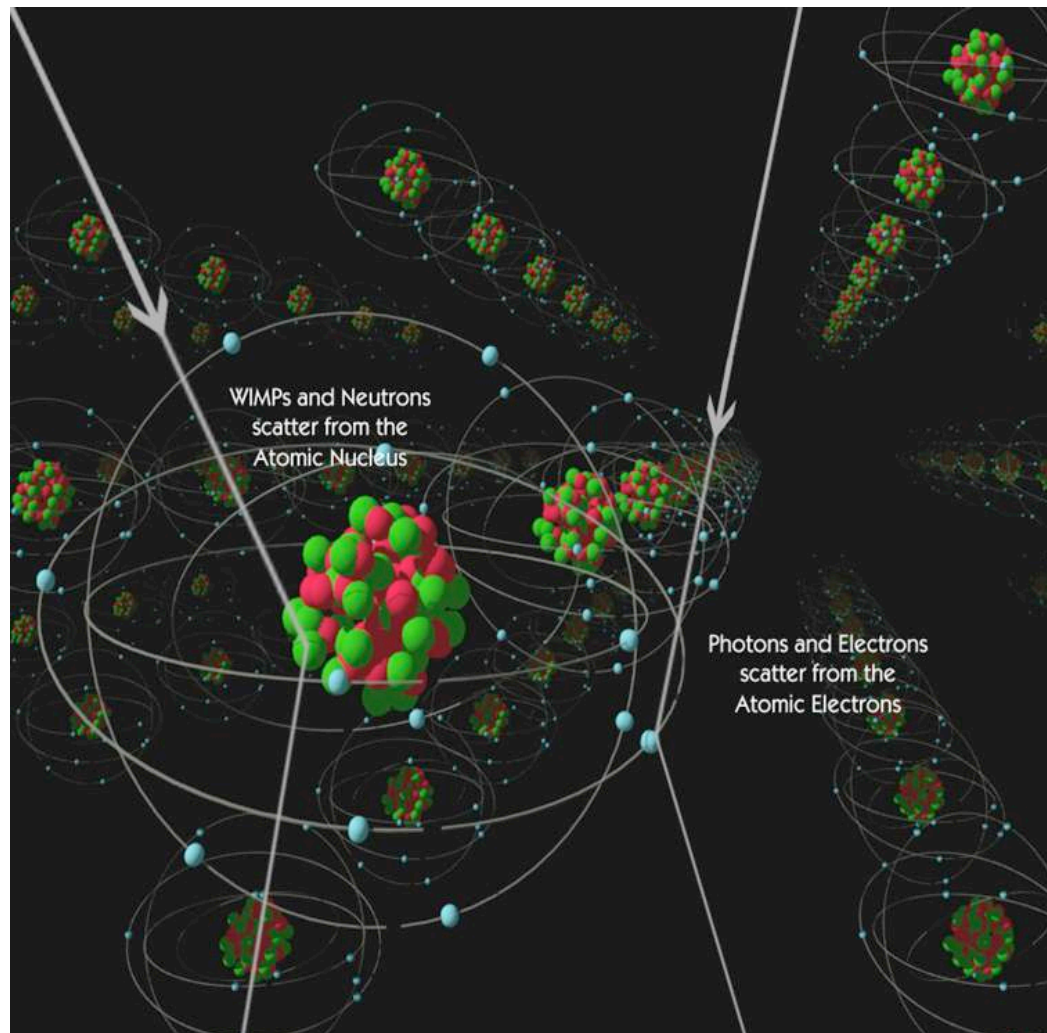
Light Dark Matter and Neutrinos from the Sun



Caroline B. Bräuninger (TUM/MPP)

*in collaboration with
Rolf Kappl and Martin W. Winkler*

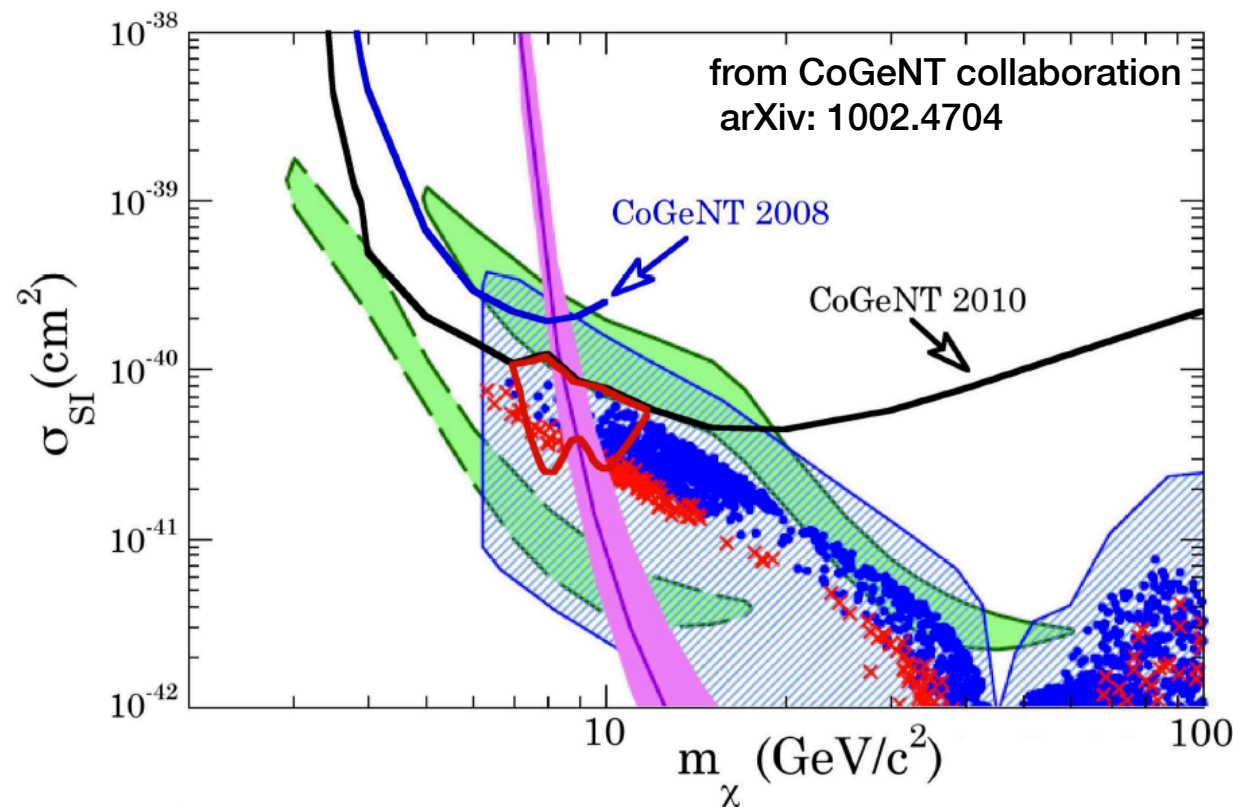
Dark Matter direct detection



Detection techniques:

- ionization
- heat
- scintillation

A direct detection of Dark Matter?



CoGeNT

- * rise in spectrum at low energies
- * background? no satisfactory explanation

CDMS

- * 2 events survive all discrimination and subtraction procedures

DAMA/Libra

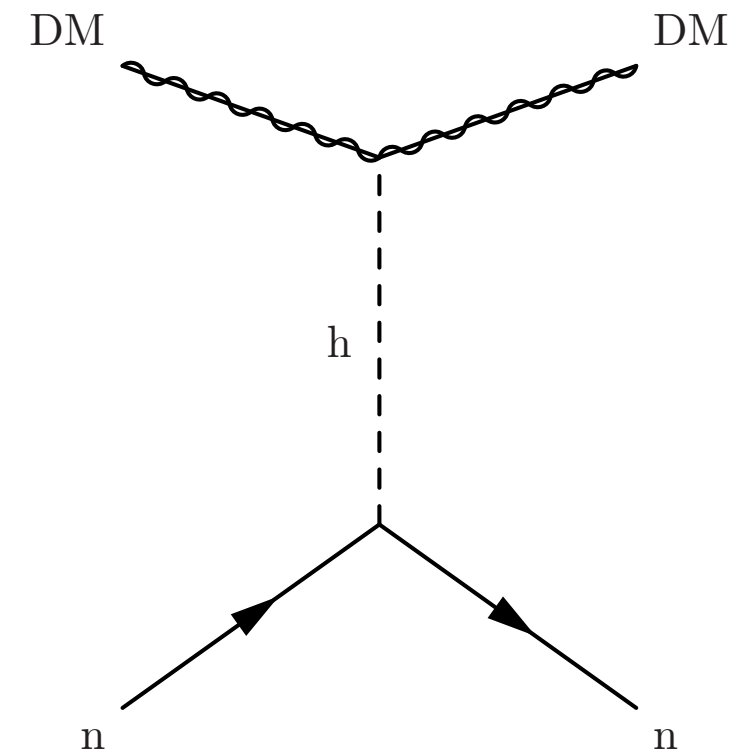
- * observe annual modulation
- * channeling effect? different mass/cross section ranges

WIMPs with $m \sim 5 - 10 \text{ GeV}$ and $\sigma_n \sim 10^{-40} \text{ cm}^2$ could fit all these signals.

Which SUSY models accomodate such WIMPs?

MSSM: Higgs known to be heavy
 \Rightarrow cross section suppressed

**Solution: singlet-extended
SUSY models allow for lighter
Higgs bosons as intermediate
particles!**



N.B.: There is also a recent paper where a complex singlet is added to SM (CSM). They can also explain these signals.

Barger, Gao, McCaskey, Shaughnessy '10

Singlet-extended SUSY models

MSSM + gauge singlet superfield S

$$W = \mu H_u H_d + \lambda S H_u H_d + \frac{\mu_s}{2} S^2 + \frac{\kappa}{3} S^3 + \text{Yukawa}$$

NMSSM: Z_3 symmetry \rightarrow solution to μ - problem

S-MSSM: allow all terms \rightarrow no solution to μ - problem

New particles: CP-even and CP-odd scalar h_s and a_s ,
singlino \tilde{s}

Decoupling limit of MSSM Higgs sector: mixing only with
light MSSM Higgs:

Two mass eigenstates:

h_1 mainly singlet

$h_2 \sim$ light MSSM Higgs

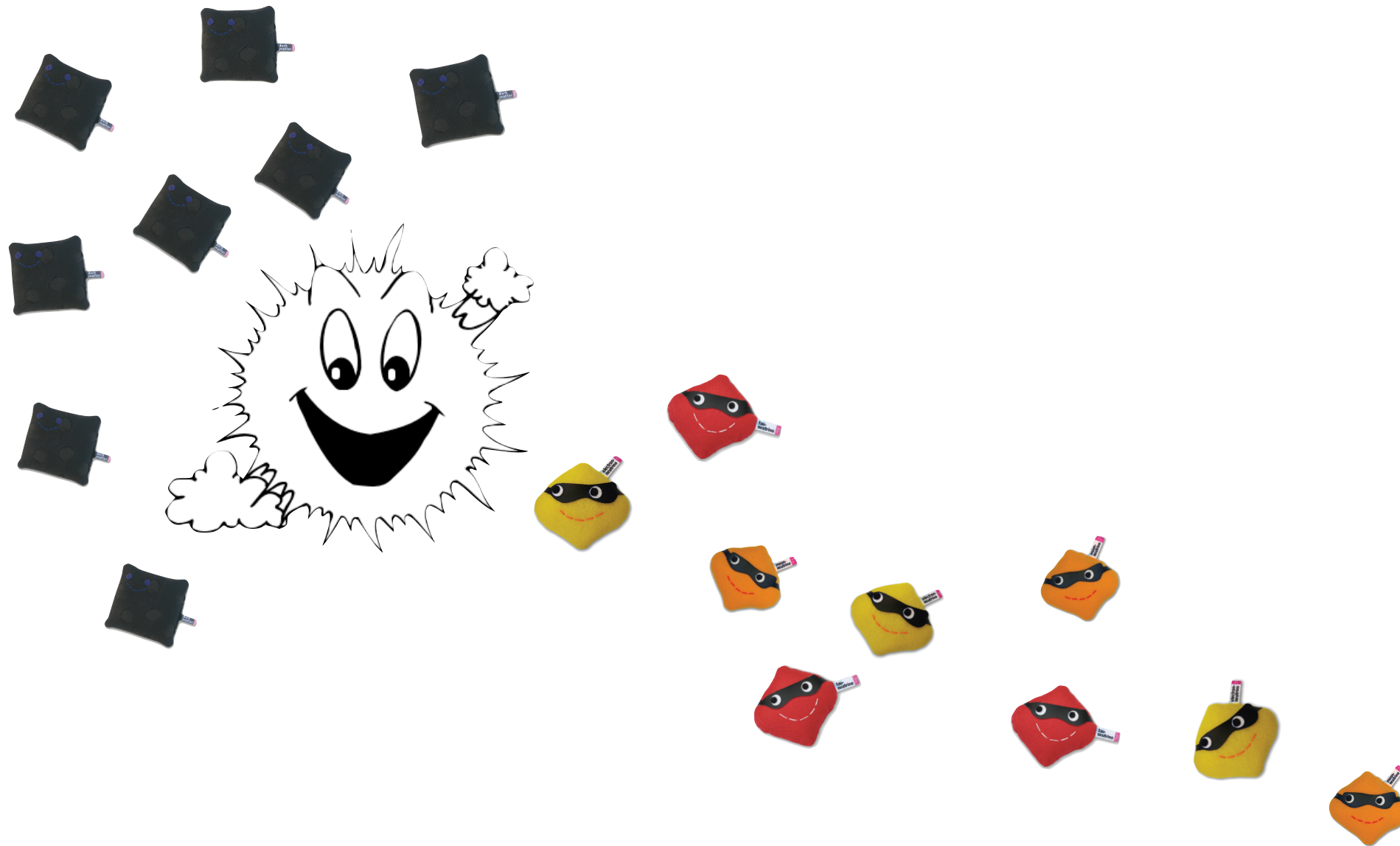
Can singlet-extended SUSY models explain CoGeNT/DAMA?

NMSSM: lightest neutralino mostly bino \rightarrow small coupling to singlet like Higgs \rightarrow cross section too small

S-MSSM: singlino DM with low mass & correct cross section thanks to mediation by light mainly singlet Higgs

Kappl, Ratz, Winkler '10
Belikov, Gunion, Hooper, Tait '10

Do we expect other signals from light singlino DM?

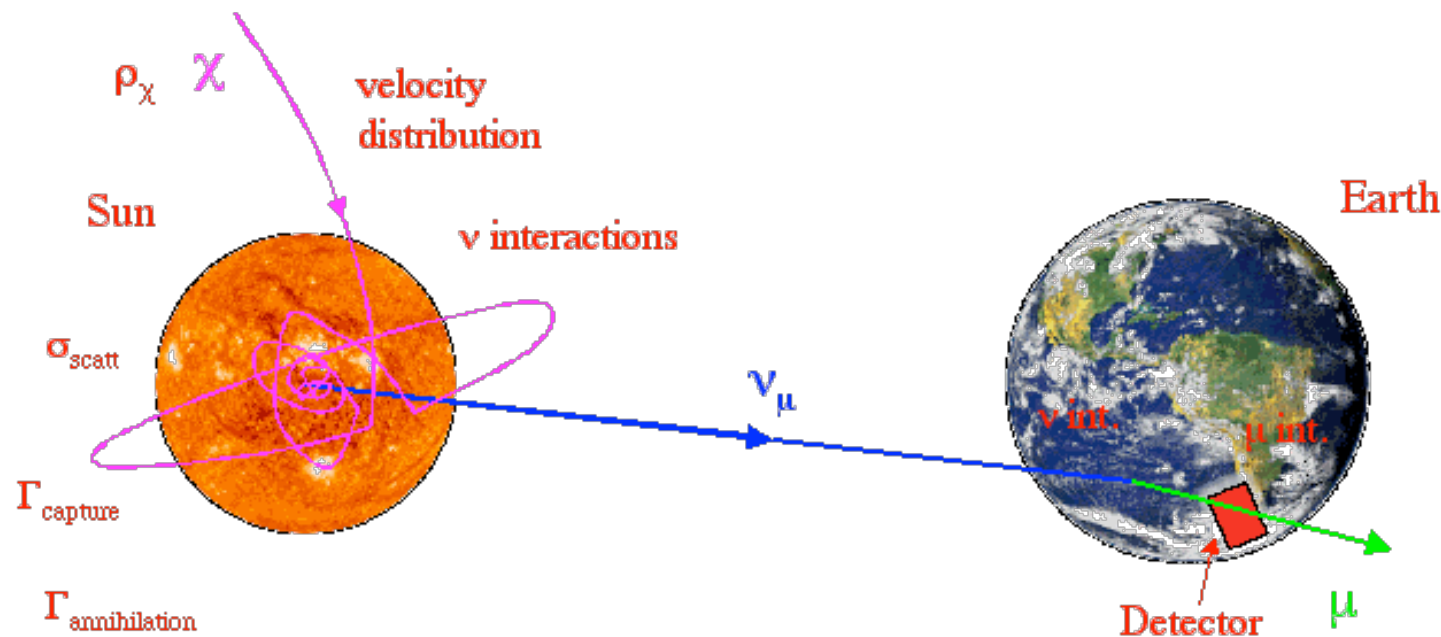


Dark Matter capture and annihilation in the Sun

DM is captured inside Sun \leftrightarrow annihilations deplete DM population

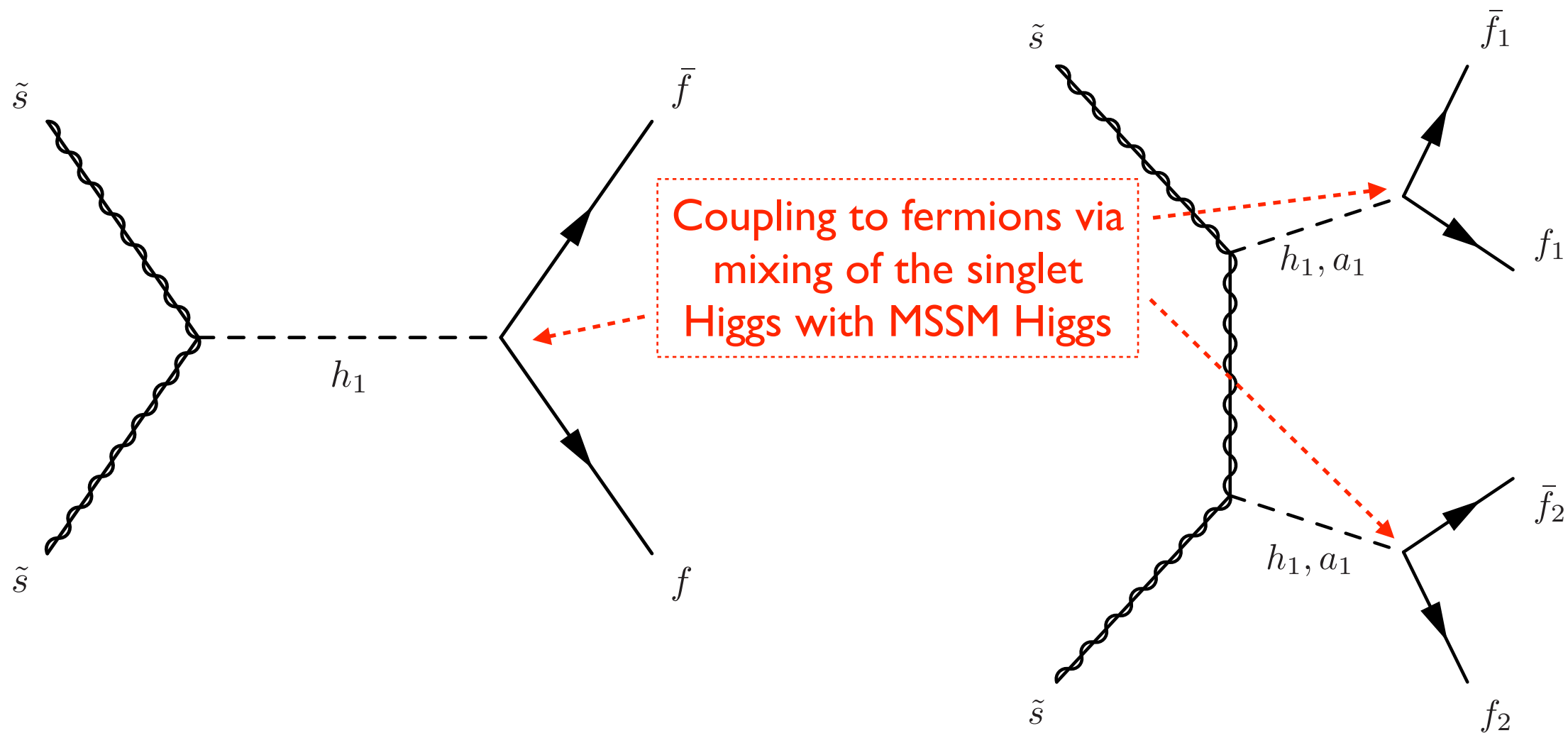
$$\Gamma_{\text{ann.}} = \frac{\Gamma_{\text{capt.}}}{2} \tanh^2(t_0/\tau_A)$$

Capture/annihilation equilibrium is reached: $\Gamma_{\text{ann.}} = \frac{\Gamma_{\text{capt.}}}{2}$



$$\Gamma_{\text{capt.}} \propto \sigma_n \rho_{\text{DM}} \left(\frac{1}{m_{\text{DM}}} \right)^2$$

Annihilations of singlino Dark Matter



- ✦ Hadronization of quarks and lepton decay simulated with Pythia
- ✦ Free decay only if $\gamma\tau_{\text{dec.}} < \tau_{\text{stop}}$

Neutrino propagation: oscillations, scatterings...



Cirelli, Fornengo, Montaruli, Sokalski, Strumia, Vissani '05
Blennow, Edsjö, Ohlsson '07

$$\frac{d\rho}{dr} = -i[H, \rho] + \left. \frac{d\rho}{dr} \right|_{\text{CC}} + \left. \frac{d\rho}{dr} \right|_{\text{NC}} + \left. \frac{d\rho}{dr} \right|_{\text{injected}}$$

oscillations

scatterings
only relevant for
 $E_\nu \gtrsim 100 \text{ GeV}$

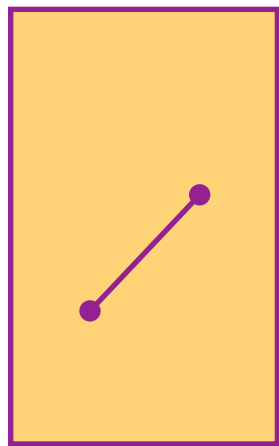
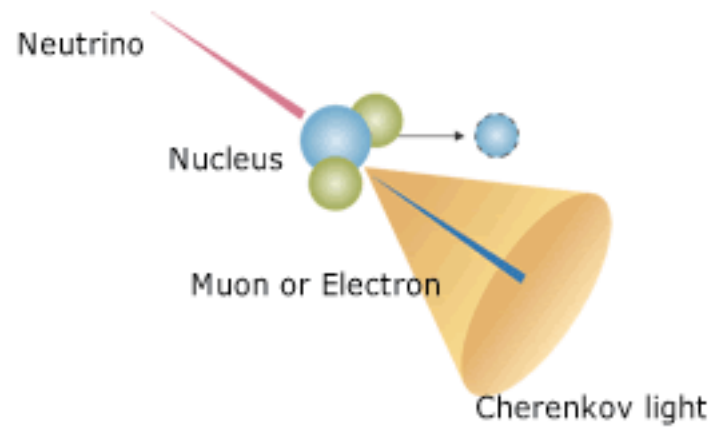
injected spectrum

$$H = \frac{m^\dagger m}{2E_\nu} + \sqrt{2}G_F N_e \text{diag.}(1, 0, 0)$$

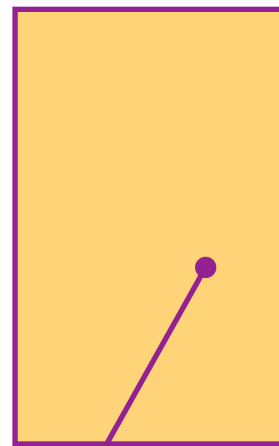
Oscillations are averaged by:

- ★ large variation in baseline
- ★ finite energy resolution of detector

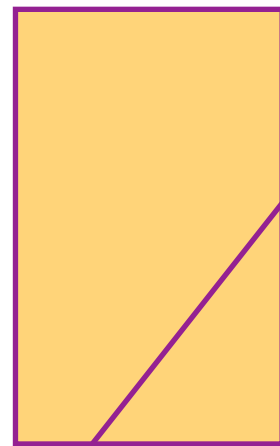
Detection of neutrinos at Superkamiokande



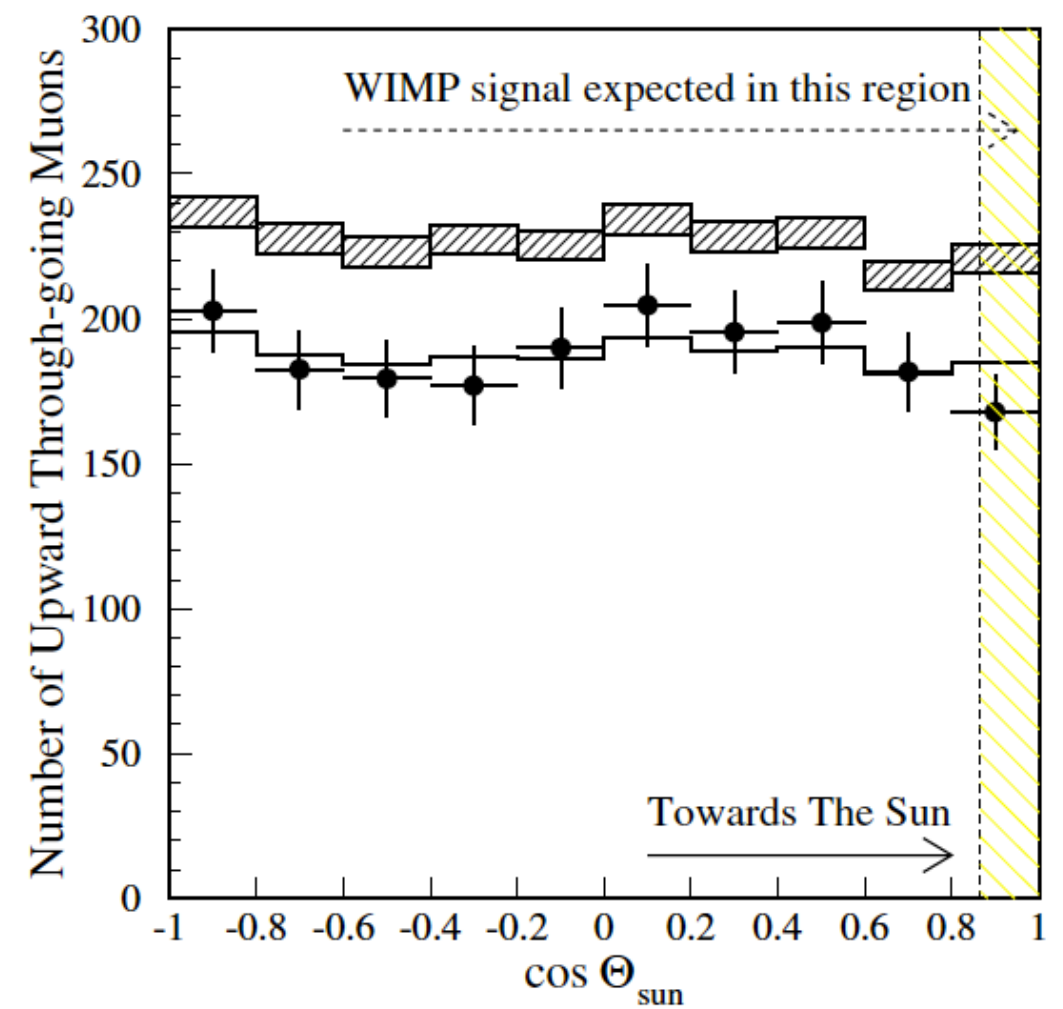
Fully-contained



Partially-contained



up-ward
through-
going



Conclusions

- Direct detection of DM? If, yes: light DM with relatively large scattering cross section
- Possible in singlet extended MSSM
- Possible signal: neutrinos generated in DM annihilations inside Sun:
 - ~ Accumulation of DM inside Sun
 - ~ Annihilation + showering
 - ~ Propagation of neutrinos to the Earth: oscillations
 - ~ Detection of neutrinos in Super-Kamiokande

Neutrino propagation for tri-bi-maximal mixing

Lehnert, Weiler '08

$$\theta_{13} = 0$$

$$H' = \begin{pmatrix} H_{2 \times 2} & 0 \\ 0 & \frac{\Delta m_{32}^2}{2E_\nu} \end{pmatrix} \quad |3\rangle_m = |3\rangle = \sum_{\alpha} U_{\alpha 3}^* |\alpha\rangle = \sin \theta_{23} |\mu\rangle + \cos \theta_{23} |\tau\rangle$$

10 MeV ~ 10 GeV

adiabatic approximation:

mass eigenstates in matter \approx instantaneous eigenstates that diagonalize $H_{2 \times 2}$ at point r , no level crossings

Centre of the Sun: $|1, 0\rangle \simeq |e\rangle$, $|2, 0\rangle = \cos \theta_{23} |\mu\rangle - \sin \theta_{23} |\tau\rangle$

$$\rho(0) = w_e |e\rangle \langle e| + w_\mu |\mu\rangle \langle \mu| + w_\tau |\tau\rangle \langle \tau|.$$

Off-diagonal elements acquire phase factor $\exp(-i\Delta m_{kj}^2 L/2E_\nu)$

→ averages to zero → drop off-diagonal elements

Evolution through the sun: $|1, 0\rangle \rightarrow |2\rangle$, $|2, 0\rangle \rightarrow |1\rangle$, $|3, 0\rangle \rightarrow |3\rangle$

$$P_{\nu_\odot \rightarrow \nu_\beta} = \langle \beta | \rho(r > R_\odot) | \beta \rangle = w_e |U_{\beta 2}|^2 + \frac{1}{2}(w_\mu + w_\tau) (1 - |U_{\beta 2}|^2)$$

Tribimaximal mixing: $P_{\nu_\odot \rightarrow \nu_e} = P_{\nu_\odot \rightarrow \nu_\mu} = P_{\nu_\odot \rightarrow \nu_\tau} = \frac{1}{3}$