25 minutes

Effective field theory for quantum fields in de Sitter space

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Andrea Federico Sanfilippo Physik-Department, Technische Universität München

with Martin Beneke, Patrick Hager

IMPRS Colloquium







Outline of the talk

- 1. Motivation: why an EFT in de Sitter space?
- 2. Soft de Sitter Effective Theory
- 3. Our work
- 4. Summary and outlook

Inflationary era in the early Universe: Quasi-exponential expansion of the Universe very shortly after Big Bang.

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- Explains approximate spatial flatness and isotropy of the Universe we see today.
- Explains approximate isotropy of the CMB and offers a source for its temperature fluctuations as well.
- Offers a mechanism for the seeding of large-scale structure.



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Source: Cherenkov Telescope Array

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with a very flat potential $V(\Phi)$. The fluctuation ϕ is well-described as a massless scalar field living in the inflationary spacetime. Background geometry is also allowed to fluctuate, have metric perturbations γ_{ij} . Main observables of interest are cosmological correlators of inflaton and metric fluctuations,

$$\langle \phi(t, \vec{x}_1) \dots \phi(t, \vec{x}_m) \gamma_{ij}(t, \vec{x}_{m+1}) \dots \gamma_{ab}(t, \vec{x}_n) \rangle$$

evaluated at equal and late time t, a long time after inflation ended (i.e. today).

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$$\mathrm{d}s^2 = \mathrm{d}t^2 - a(t)^2 \mathrm{d}\vec{x}^2 \,, \quad a(t) \equiv e^{Ht} \,,$$

with H = const. the Hubble parameter. Observer sees a cosmological horizon of radius r = 1/H.

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► Massless scalar propagator is infrared-divergent ⇒ cannot construct perturbative expansion:

 \Rightarrow Perturbation theory is ill-defined, need to go beyond it to obtain physically sensible results.

 IR-regularized perturbation theory (e.g. with small mass m² « H² or IR-cutoff) develops secular logarithms:

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Even for weak coupling g the perturbative expansion of an observable O breaks down after sufficiently long time, e.g.

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 \Rightarrow In both cases, perturbation theory fails due to field modes with

$$\frac{k}{a(t)H} \ll 1 \,.$$

The IR-structure of this theory appears to be non-trivial and needs to be understood in a framework that goes beyond fixed-order perturbation theory.

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$$\begin{split} \phi(t,\vec{k}) &= \bar{\phi}(t,\vec{k}) + \delta\phi(t,\vec{k}) \,, \quad \delta\phi \ll \bar{\phi} \,, \\ \uparrow & \uparrow \\ k < a(t)H \quad k > a(t)H \end{split}$$

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The stochastic "noise" is generated by the short-wavelenegth modes $\delta\phi$.

The PDF determines all one-point correlation functions of $\bar{\phi}$ via

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under the assumption of an equilibrium solution, $\partial_t \rho = 0$, we find

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 \Rightarrow Interpretation: the massless, interacting field develops a dynamical mass

$$m_{\rm dyn}^2 \sim H^2 \sqrt{g} \,,$$

non-perturbatively. The IR-divergence self-regularizes.

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 \Rightarrow This motivates the construction of an effective framework in the context of QFT to understand how to systematically incorporate subleading corrections.

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- Exhibit the property of "factorization" of hard and soft physics, schematically:

 $\lim_{-k_i/(a(t)H)\to 0} \langle \phi(t,\vec{k}_1)...\phi(t,\vec{k}_n)\rangle = C_{\rm hard} \times \langle \varphi(t,\vec{k}_1)...\varphi(t,\vec{k}_n)\rangle_{\rm EFT}\,.$

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Breakthrough in this direction: Soft de Sitter Effective Theory (SdSET) [Cohen, Green 2020; Cohen, Green, Premkumar, Ridgway 2021; Cohen, Green, Premkumar 2021].

The SdSET is an EFT built to compute the late-time/small-wavenumber limit of equal-time correlation functions,

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Full-theory field is split up as

$$\phi(t,\vec{k}) = \underbrace{H\left[\varphi_{+}(t,\vec{k}) + [a(t)H]^{-3}\varphi_{-}(t,\vec{k})\right]}_{\text{EFT fields, } k/\Lambda(t) < 1} + \phi_{\text{UV}}(t,\vec{k}),$$

where for fixed t the quantity

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 $\phi_{\rm UV}$ is not described as a dynamical field in the EFT (it is "integrated out"). Its effects are captured by Wilson coefficients and non-Gaussian initial conditions (IC's).





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Physical origin of the IC's:



SdSET describes the time evolution from t_H on, field modes already had time to evolve and interact \Rightarrow leads to non-Gaussian correlations at t_H , which need to be specified as input for the EFT via the IC's.

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In a bit more detail: SdSET is defined at leading power in λ by the action

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► The Ξ_n are time-independent, non-local functions which encode the non-Gaussian IC's, they are treated as non-local "vertices" in computations. The IC-sector of the action is localized at the time of horizon-crossing t_H.

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Both the Ξ_n and c_n are determined by matching computations: compute quantity in the full theory, compute analogous quantity in the EFT, the difference determines Ξ_n and c_n .

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Since in dS there is no consensus on a good regulator for divergent loop integrals, need to define regularization and renormalization schemes both in the full theory and EFT. Introduced an analytic regulator via

$$\phi(t,\vec{x}) \equiv \int \frac{\mathrm{d}^3 k}{(2\pi)^3} e^{i\vec{k}\cdot\vec{x}} \left(\frac{k}{a(t)\mu}\right)^{-\varepsilon} \phi(t,\vec{k}) \,,$$

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• Opportunity to apply tools from flat-space computations to cosmological correlators. In particular: applied the method of regions [Beneke, Smirnov 1997] to simplify computation of correlators in the late-time limit and elucidate the origin of the contributions to the Ξ_n and c_n .

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$$\phi \longrightarrow \phi + \phi \longrightarrow \phi \stackrel{!}{=} \varphi_{+} \longrightarrow \varphi_{+} + \varphi_{+} \longrightarrow \varphi_{+}$$

 Ξ_4 , c_4 already determined, need to be able to reproduce full result only using Ξ_2 , $c_2 \Rightarrow$ first consistency check \checkmark .

Finally, matching of composite-operator correlation function

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 \Rightarrow Would like to derive this rigorously from an RGE analysis and put the framework on a solid QFT footing.

Summary



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- To simplify full-theory calculations and elucidate the origin of the contributions to Wilson coefficients and IC's we were able to use the method of regions.

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Thank you for your attention!

Backup

Approximately exponential inflation of spacetime occurs if

$$V(\phi_0) \gg \left(\frac{\partial \phi_0}{\partial t}\right)^2.$$

 $\Rightarrow \phi_0$ "slowly rolls down" a flat potential V.

